

# THE EVOLUTION OF ACTIVE GALACTIC NUCLEI IN CLUSTERS OF GALAXIES TO REDSHIFT 1.3

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## ABSTRACT

We have measured the luminous AGN population in a large sample of clusters of galaxies and find evidence for a substantial increase in the cluster AGN population from  $z \sim 0.05$  to  $z \sim 1.3$ . The present sample now includes 32 clusters of galaxies, including 15 clusters above  $z = 0.4$ , which corresponds to a three-fold increase compared to our previous work at high redshift. At  $z < 0.4$  we have obtained new observations of AGN candidates in six additional clusters and found no new luminous AGN in cluster members. Our total sample of 17 low-redshift clusters contains only two luminous AGN, while at high redshifts there are 18 such AGN, or an average of more than one per cluster. We have characterized the evolution of luminous X-ray AGN as the fraction of galaxies with  $M_R < M_R^*(z) + 1$  that host AGN with rest-frame, hard X-ray [2–10 keV] luminosities  $L_{X,H} \geq 10^{43} \text{ erg s}^{-1}$ . The AGN fraction increases from  $f_A = 0.134^{+0.18}_{-0.087} \%$  at a median  $z = 0.19$  to  $f_A = 1.00^{+0.29}_{-0.23} \%$  at a median  $z = 0.72$ . Our best estimate of the evolution is a factor of eight increase to  $z = 1$  and the statistical significance of the increase is  $3.8\sigma$ . This dramatic evolution is qualitatively similar to the evolution of the star-forming galaxy population in clusters known as the Butcher-Oemler effect. We discuss the implications of this result for the coevolution of black holes and galaxies in clusters, the evolution of AGN feedback, searches for clusters with the Sunyaev-Zel'dovich effect, and the possible detection of environment-dependent downsizing.

*Subject headings:* galaxies: active – galaxies: clusters: general – galaxies: evolution – X-rays: galaxies – X-rays: galaxies: clusters – X-rays: general

## 1. INTRODUCTION

The demographics of AGN in clusters of galaxies have important implications for the growth of the supermassive black holes at the centers of cluster galaxies, the nature of AGN fueling, and the impact of AGN on the intracluster medium (ICM) over cosmic time. The luminous, massive elliptical galaxies that dominate the galaxy population in the richest clusters are also expected (and in some cases are measured: Houghton et al. 2006; Gebhardt et al. 2007) to have the most massive black holes in the local universe. As the stars in these galaxies appear to have an earlier mean formation epoch than those in field galaxies (e.g. van Dokkum & Franx 1996; Kelson et al. 1997), the apparent coevolution of black holes and galaxies (e.g. Hopkins et al. 2006, and references therein) implies that the bulk of their present black hole mass was also accreted at earlier times.

This scenario is also motivated by observations of local clusters that clearly show their galaxy populations are more quiescent than local field galaxies. An early demonstration by Osterbrock (1960) showed that cluster ellipticals were far less likely to have [O II]  $\lambda 3727$  emission than field ellipticals, a result that has since been confirmed by many studies (e.g. Gisler 1978; Dressler et al. 1985, 1999). One big question that has motivated this work is: Why are galaxy populations different in clusters? Numerous physical mechanisms have been invoked to explain the relative lack of star formation in cluster galaxies, as well as their higher fraction of elliptical and S0 galaxies (Dressler 1980) and relative lack of cold gas (e.g. Giovanelli & Haynes 1985). These include ram-

pressure stripping by the ICM (Gunn & Gott 1972), evaporation of a galaxy's interstellar medium (ISM) by the hot ICM (Cowie & Songaila 1977), tidal effects with the cluster potential (Farouki & Shapiro 1981; Merritt 1983; Byrd & Valtonen 1990), the absence of newly-accreted cold gas (Larson et al. 1980), and galaxy harassment and mergers (Richstone 1976; Moore et al. 1996).

All of these physical effects may also be important for fueling accretion onto the central black holes in galaxies because they impact either the available gas supply in a galaxy, angular momentum transport, or both. The best and perhaps only candidate process for fueling the most luminous AGN is the merger of two gas-rich galaxies (e.g. Barnes & Hernquist 1992) and the relative lack of both cold gas and major mergers is a reasonable explanation for the nearly complete absence of QSOs hosted by cluster galaxies. For less luminous AGN the case is less clear because an increasing number of physical processes such as minor mergers, galaxy harassment, various types of bars, stellar mass loss, etc. could also play a role (see Martini 2004, for a review). If mechanisms such as galaxy harassment and stellar mass loss are important for fueling low-luminosity AGN, then comparable numbers of low-luminosity AGN may be present in clusters and the field.

Recent studies of the AGN fraction as a function of environment with emission-line galaxies from the Sloan Digital Sky Survey (SDSS) find that the most luminous AGN are rarer in denser environments (SDSS; Kauffmann et al. 2004; Popesso & Biviano 2006), although these studies do not sample the densest regions of clusters well. This decrease is in contrast to both lower-luminosity AGN in SDSS (Miller et al. 2003) and radio observations (Best 2004; Best et al. 2005), which show that the radio AGN fraction does not decrease

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significantly in denser environments. X-ray observations with *Chandra* show that the X-ray AGN fraction is larger than expected from AGN selection via visible-wavelength emission-lines. In previous work we showed that X-ray observations identified approximately five times as many AGN as selection at visible-wavelengths (Martini et al. 2002, 2006), although the precise value of the X-ray excess depends significantly on the relative sensitivity and luminosity threshold of the observations. This spectroscopic study of X-ray counterparts confirmed the many previous studies that suggested a higher X-ray AGN population in clusters from surface density arguments alone (e.g. Cappi et al. 2001; Sun & Murray 2002; Ruderman & Ebeling 2005), yet it is still not clear if the X-ray AGN fraction is higher than the field value. To date there is only weak evidence that the X-ray AGN fraction in clusters is comparable to the fraction in field early-type galaxies (Lehmer et al. 2007; Sivakoff et al. 2008; Arnold et al. 2009). One of the virtues of the emission-line galaxy studies as a function of environment is that they can directly calculate the fraction of a given galaxy population that hosts AGN as a function of environment, even though this technique appears to systematically miss AGN in the densest regions relative to X-ray and radio selection.

In addition to a local comparison between AGN in different environments, measurement of the evolution of the AGN population in clusters can constrain the formation epoch for their supermassive black holes and the extent of their coevolution with the cluster galaxy population. The key early work on the evolution of galaxies in clusters was by Butcher & Oemler (1978, 1984), who observed a substantial increase in the fraction of blue galaxies in higher-redshift clusters. The Butcher-Oemler effect is interpreted as an increase in the amount of star formation and has been confirmed by many other indicators, in particular [O II] emission-line galaxy fractions (Poggianti et al. 2006) and an increase in the number of  $24\mu\text{m}$  sources in *Spitzer* observations of distant clusters (Bai et al. 2007; Saintonge et al. 2008). The observed increase brings the star formation rate (SFR) in cluster galaxies closer to those in the field. At a redshift of  $z \sim 1$  and higher, observations with *Spitzer* even find that galaxies in denser environments have higher star formation rates than lower-density regions (Elbaz et al. 2007), which is opposite the trend observed in the local universe. Similar results have also been found with deep UV data (Heinis et al. 2007). The situation is less clear when star formation is measured with the [O II] emission line because while Poggianti et al. (2008) find that star formation does not strongly depend on environment, Cooper et al. (2008) find the specific star formation rate has a similar dependence on environment at  $z = 0$  and  $z = 1$ , although the total star formation rate is higher in clusters at  $z = 1$  than in the field.

The existence of the Butcher-Oemler effect and the many indirect arguments outlined above for a connection between star formation and black hole accretion suggest that there should be an increased AGN population in high-redshift clusters. An early study of the high-redshift cluster 3C295 at  $z = 0.46$  by Dressler & Gunn (1983) found evidence for three AGN and was an indication that this may be the case; however, their relative scarcity precluded a detailed statistical study or targeted studies to deliberately identify cluster AGN. This situation changed dramatically with the launch of *Chandra*, whose superb sensitivity and angular resolution produced a dramatic increase in efficiency for searches for AGN, particularly lower-luminosity sources. Just as the

case for local clusters, *Chandra* observations of distant clusters have revealed substantial populations of point sources (Cappelluti et al. 2005; Gilmour et al. 2009). Spectroscopic confirmation that these point sources are associated with cluster members has been more challenging (Johnson et al. 2003; Demarco et al. 2005), but in Eastman et al. (2007) we combined new observations of MS2053.7-0449 ( $z = 0.58$ ) with archival data on three additional,  $z > 0.5$  clusters and found an approximately order of magnitude increase in the fraction of  $M_R < -20$  mag galaxies that hosted AGN more luminous than  $L_{X,H} \geq 10^{43} \text{ erg s}^{-1}$  in the hard X-ray band (2–10 keV) relative to the sample of ten low-redshift  $z < 0.32$  clusters in Martini et al. (2007). These results have since been strengthened with detailed studies of clusters at  $z \sim 1$  with XMM (van Breukelen et al. 2009) and measurements of surface density excesses in clusters to  $z \sim 1.5$  (Galametz et al. 2009).

In addition to their application to the coevolution of black holes and galaxies, an increase in the AGN fraction in clusters may also impact the ICM. At low redshifts many studies have shown that AGN feedback is a viable explanation for the absence of substantial reservoirs of cold gas at the centers of clusters (for a recent review see McNamara & Nulsen 2007). This feedback is ascribed to AGN associated with the central cluster galaxy, which is almost invariably a luminous radio source. In our studies of X-ray AGN this is almost the only cluster galaxy in which we are *insensitive* to the presence of an AGN because it is challenging to measure even a bright nuclear point source when juxtaposed with the extended emission from the ICM that often peaks near the central cluster galaxy. Nevertheless, the evolution of AGN in other cluster galaxies is likely to be connected to the evolution of the central AGN as the stars in the most luminous cluster galaxies have comparable ages. An increase in the net energy production by AGN in higher-redshift clusters is of interest because energy input during cluster formation has been invoked as an explanation for the minimum entropy level in the ICM (Kaiser 1991; Evrard & Henry 1991). AGN remain perhaps the most viable mechanism, if only because most others can be ruled out (Kravtsov et al. 2004; Conroy & Ostriker 2008), although the details of how AGN feedback couples to the ICM remain uncertain. Outside of the central galaxy, an increase in the number of other AGN associated with clusters of galaxies may also affect measurement of other cluster properties (Branchesi et al. 2007; Bignamini et al. 2008). Finally, an analogous increase in the radio-loud AGN population in high-redshift clusters may contaminate searches for clusters via the Sunyaev-Zel'dovich effect (Sunyaev & Zeldovich 1970) at mm and cm wavelengths. As many searches for clusters that exploit this effect are in progress, it is important to characterize the potential impact of evolution of the cluster AGN population on these experiments (e.g. Lin & Mohr 2007).

In the next section we describe our expanded high-redshift data, as well as the selection criteria for X-ray AGN we employ at all redshifts. We then describe our new observations of low-redshift clusters in §3. These two datasets are combined to calculate the cluster AGN fraction and its evolution in §4, followed by an examination of the properties of the cluster AGN in §5. We discuss the implications of these results, particularly on the coevolution of black holes and galaxies, in §6 and conclude with a summary of our results. Throughout this paper we assume that the cosmological parameters are:  $(\Omega_M, \Omega_\Lambda, h) = (0.3, 0.7, 0.7)$  where  $H_0 = 100h \text{ km s}^{-1} \text{ Mpc}^{-1}$ . All absolute magnitudes quoted in this paper assume  $h = 0.7$ .

## 2. HIGH-REDSHIFT DATA

Two large surveys have obtained redshifts for substantial numbers of galaxies with X-ray counterparts in many deep, archival *Chandra* observations that include substantial numbers of high-redshift clusters of galaxies. These are the Serendipitous Extragalactic X-ray Source Identification Program (SEXSI; Harrison et al. 2003; Eckart et al. 2005, 2006) and the Chandra Multiwavelength Project (ChAMP; Kim et al. 2004a,b; Green et al. 2004; Silverman et al. 2005). We have investigated the fields surveyed by both SEXSI and ChAMP to identify datasets that contain clusters of galaxies with  $z > 0.4$  and have sufficient depth to identify  $L_{X,H} \geq 10^{43}$  erg s $^{-1}$  (rest frame 2–10 keV) AGN at the cluster redshift.

The SEXSI survey published spectroscopic redshifts for 27 archival *Chandra* observations in Eckart et al. (2006) that were selected to identify hard X-ray sources over the flux range of  $10^{-13} - 10^{-15}$  erg s $^{-1}$ cm $^{-2}$  and isolate those responsible for the hard X-ray background. The specific selection criteria for the fields were that they must be high Galactic latitude ( $|b| > 20^\circ$ ) and be obtained with either the I or S modes of the Advanced Camera for Imaging Spectroscopy (ACIS; Bautz et al. 1998) when no grating was used. The X-ray luminosities quoted by SEXSI are based on spectral fits that assume a  $\Gamma = 1.5$  power law and intrinsic absorption  $N_H$  at the source redshift, although they quote the observed luminosities (not corrected for obscuration) and provide the best-fit  $N_H$  value. The average spectroscopic completeness is 67% (see §4.2 below) for sources with  $R < 24.4$  mag on the Vega system. Nine of the 27 SEXSI fields include clusters of galaxies with  $z > 0.4$  and we include seven<sup>2</sup> in our sample. As one field contains 3 clusters, we list nine clusters from SEXSI in Table 1.

The ChAMP survey published spectroscopic redshifts for 20 archival *Chandra* observations in Silverman et al. (2005) that were similarly selected for depth, high Galactic latitude ( $|b| > 20^\circ$ ), and no special observing modes. The spectroscopic completeness of ChAMP is 77% at  $r' < 22.5$  mag, where  $r'$  is on the SDSS photometric system (Fukugita et al. 1996, and  $r'_{AB} = R_{Vega} + 0.17$ ). Their X-ray luminosities are based on spectral fits that assume a  $\Gamma = 1.9$  power law and intrinsic absorption  $N_H$  at the source redshift, as well as the appropriate Galactic absorption, although they also quote the observed luminosities (only corrected for Galactic absorption). The final sample presented in Silverman et al. (2005) was restricted to X-ray sources with  $L_X > 10^{42}$  keV in the 2–8 keV band in order to insure all are AGN. Most (69%) are spectroscopically classified as broad-line AGN (BLAGN). Nine of these 20 ChAMP fields include clusters of galaxies with  $z > 0.4$  and we include eight<sup>3</sup> of these in our study (see Table 1). Two of these clusters are common to both ChAMP and SEXSI (MS2053.7-0449 and RXJ1716.4+6708) and therefore the final sample has fifteen clusters with  $z > 0.4$ . While spectroscopic data for X-ray sources in other high-redshift clusters exist (e.g. Johnson et al. 2006), we limit our high-redshift sample to these fifteen to maximize the uniformity of the dataset.

<sup>2</sup> RX J1350.0+6007 was not targeted for spectroscopy and the X-ray data for CL0442+0202 ( $z = 1.11$ ) were sufficiently shallow ( $t = 44$  ks) that they may not be complete to  $L_{X,H} = 10^{43}$  erg s $^{-1}$ . In addition, Stern et al. (2003) classify CL0442+0202 as an overdensity that has not yet collapsed, rather than as a cluster.

<sup>3</sup> We exclude CL J0152.7-1357 ( $z = 0.831$ ) because the exposure time is shorter than the others at  $t = 34.6$  ks and therefore the X-ray data may not be complete to  $L_{X,H} = 10^{43}$  erg s $^{-1}$ .

We have also compiled additional data for each cluster listed in Table 1 that will be important for our subsequent analysis. One quantity is the center of the cluster, which is needed to determine if a given AGN falls within the projected virial radius of the cluster. We associate the center of each cluster with the centroid of the extended X-ray emission. While these coordinates do not always agree with the standard coordinates quoted in the literature, this assumption makes our analysis more uniform. The redshift and velocity dispersion are also needed to determine if an AGN is within the cluster. In most cases velocity dispersions for these clusters are available in the literature and we quote the origin of the measurement we adopt in the table. When the velocity dispersion has not been measured, we estimate this quantity from the X-ray temperature and the  $\sigma - T_X$  relationship from Xue & Wu (2000). Specifically, we used the relation  $\sigma = 10^{2.51 \pm 0.01} T^{0.61 \pm 0.01}$  km s $^{-1}$  derived from their combined group and cluster sample with orthogonal distance regression (Feigelson & Babu 1992). Based on their data, we estimate that there is a 30% uncertainty in  $\sigma$  at fixed  $T$ .

One potential concern for our subsequent analysis is that the Xue & Wu (2000)  $\sigma - T$  relation may not hold at higher redshift. Lubin et al. (2004) investigated this point for several optically-selected clusters and found that they were 2-9 times cooler than expected from the local relation; however, the difference was much less stark for X-ray selected, high-redshift clusters similar (and in several cases identical to) those presented here. Fang et al. (2007) showed that high-redshift, X-ray selected clusters are consistent with the low-redshift  $L_X - \sigma$  relation, although spectroscopically-selected groups and clusters do not agree as well (see also Andreon et al. 2008).

Finally, we have calculated the projected size of the virial radius for each cluster following Treu et al. (2003) and throughout this paper we associate the virial radius with  $R_{200}$ , the radius within which the cluster is a factor of 200 overdensity. Of the three clusters we have in common with Poggianti et al. (2006), for 3C 295 and MS1054-03 we adopt nearly the same  $\sigma$  and our  $R_{200}$  estimate is nearly identical to theirs, while for MS0015.9+1609 we adopt a slightly larger velocity dispersion (1234 km s $^{-1}$  from Carlberg et al. (1996) rather than their 984 km s $^{-1}$ ) and consequently infer a larger radius.

Because the most relevant ChAMP measurements are the 2–8 keV luminosity, rather than 2–10 keV luminosity, we multiply the ChAMP 2–8 keV luminosities by a factor of 1.2. This correction factor was calculated for a  $\Gamma = 1.7$  power law with PIMMS. There is some uncertainty in this correction factor because not all AGN have this power-law form, particularly as we assume this correction for their observed rather than intrinsic (unobscured) spectra, but this is not a significant effect compared to other sources of systematic errors that we discuss below. There are no additional AGN from ChAMP that enter the sample after this step because there are none just below the  $10^{43}$  erg s $^{-1}$  threshold in the 2–8 keV band. We also estimated the difference in luminosity for an AGN calculated with the  $\Gamma = 1.5$  power law employed by SEXSI, the  $\Gamma = 1.9$  employed by ChAMP, and a  $\Gamma = 1.7$  power law to determine if these differences would cause any sources to fall in or out of the same and none would do so. In the two clusters observed by both ChAMP and SEXSI, there is one cluster AGN common to both surveys: CXOSEXSI J171636.9+670829. The redshifts from the two surveys agree exactly ( $z = 0.795$ ) and the luminosities agree well:  $L_{X,2-10} = 10^{44}$  erg s $^{-1}$  and  $L_{X,2-8} = 10^{43.88}$  erg s $^{-1}$ .

TABLE 1  
HIGH-REDSHIFT CLUSTER SAMPLE

Cluster (1)	$\alpha_c$ (2)	$\delta_c$ (3)	$z$ (4)	$\sigma$ [km/s] (5)	$\sigma$ Ref (6)	$T_X$ [keV] (7)	$T_X$ Ref (8)	$R_{200}$ [Mpc] (9)	Spectra (10)
MS 1621.5+2640	16:23:34.9	+26:34:21	0.43	735	1	7.6	1	1.42	SESXI
3C 295	14:11:20.5	+52:12:09	0.46	1642	1	5.3	1	3.12	SESXI
MS 0451.6-0305	4:54:11.1	-03:00:55	0.538	1371	2	8.1	1	2.49	ChaMP
MS 0015.9+1609	0:18:33.5	+16:26:06	0.541	1234	2	9.4	2	2.24	ChaMP
RX J0848.7+4456	8:48:47.6	+44:56:16	0.574	670	3	3.2	3	1.19	SESXI
MS 2053.7-0449	20:56:21.3	-04:37:49	0.583	865	4	5.2	1	1.53	SESXI, ChaMP
RX J0542.8-4100	5:42:49.8	-41:00:07	0.634	1101	3	7.9	3	1.89	ChaMP
RX J2302.8+0844	23:02:48.3	+08:43:48	0.722	993	3	6.6	3	1.61	ChaMP
MS 1137.5+6625	11:40:22.1	+66:08:14	0.782	967	3	6.3	1	1.52	ChaMP
RX J1317.4+2911	13:17:22.0	+29:11:24	0.805	531	3	2.2	1	0.82	SESXI
RX J1716.4+6708	17:16:49.3	+67:08:25	0.813	1445	1	6.6	1	2.22	SESXI, ChaMP
MS 1054-03	10:56:55.7	-03:37:39	0.831	1156	5	7.8	1	1.76	ChaMP
RDCS J0910+5422	9:10:44.7	+54:22:04	1.11	675	6	3.5	1	0.87	SESXI
Lynx E	8:48:58.3	+44:51:51	1.261	740	7	3.8	4	0.88	SESXI
Lynx W	8:48:34.2	+44:53:35	1.27	650	8	1.7	4	0.77	SESXI

NOTE. — Cluster sample and properties derived from the present study. Columns are: (1) Cluster name; (2 and 3) RA and DEC for the centroid of the extended X-ray emission; (4) redshift; (5) velocity dispersion; (6) reference for the velocity dispersion; (7) X-ray temperature in keV; (8) reference for the X-ray temperature; (9) estimate of the virial radius in Mpc (e.g., Treu et al. 2003); (10) origin of most of the spectra. References for velocity dispersion are: 1: Girardi & Mezzetti (2001); 2: Carlberg et al. (1996); 3: derived from the X-ray temperature following Xue & Wu (2000); 4: Tran et al. (2005); 5: Tran et al. (2007); 6: Mei et al. (2006); 7: from weak lensing estimate Jee et al. (2006); 8: Stanford et al. (2001). References for X-ray temperatures are: 1: Vikhlinin et al. (2002); 2: Ebeling et al. (2007); 3: Ettori et al. (2004); 4: Jee et al. (2006).

TABLE 2  
HIGH-REDSHIFT CLUSTER AGN SAMPLE

AGN (1)	Cluster (2)	$z$ (3)	$R$ [mag] (4)	$\log L_{X,H}$ [erg s <sup>-1</sup> ] (5)	$\delta v/\sigma$ (6)	$\Delta R$ [arcmin] (7)	$R/R_{200}$ (8)	Class (9)
CXOSEXSI J141127.4+521131	3C295	0.451	19.78	43.4	1.13	1.23	0.14	ALG
CXOSEXSI J141123.4+521331	3C295	0.472	19.05	43.8	1.5	1.45	0.16	BLAGN
E0015+162	MS0015.9+1609	0.553	18.41	45.48	1.89	3.35	0.58	BLAGN
CXOSEXSI J084858.0+445434	RX J0848.7+4456	0.573	19.58	43.8	0.28	2.5	0.83	BLAGN
CXOMP J054248.2-410140	RDCSJ0542-4100	0.634	20.64	43.24	0	1.58	0.32	NELG
CXOMP J054251.4-410205	RDCSJ0542-4100	0.637	19.63	43.35	0.5	1.99	0.33	ALG
CXOMP J054259.5-410241	RDCSJ0542-4100	0.638	20.50	43.37	0.67	3.16	0.63	NELG
CXOMP J054240.8-405626	RDCSJ0542-4100	0.639	20.89	43.67	0.83	4.05	0.81	NELG
CXOMP J054255.0-405922	RDCSJ0542-4100	0.644	22.08	43.08	1.67	1.24	0.25	NELG
CXOMP J114022.0+660816	MS1137+6625	0.786	20.37	43.24	0.7	0.04	0.01	BLAGN
CXOSEXSI J171636.9+670829	RXJ1716.4+6708	0.795	22	44	2.06	1.19	0.24	ELG
CXOSEXSI J131718.8+291111	RX J1317.4+2911	0.803	21.98	43.3	0.63	0.68	0.38	BLAGN
CXOSEXSI J171703.8+670900	RXJ1716.4+6708	0.812	21.79	43	0.11	1.53	0.31	ELG
CXOSEXSI J171714.5+671136	RXJ1716.4+6708	0.815	22.68	43.2	0.23	4.02	0.82	ELG
CXOMP J105650.6-033508	MS 1054-03	0.818	21.76	43.22	1.84	2.82	0.73	BLAGN
CXOU J091043.3+542152	RDCSJ0910+5422	1.104	24	43.06	1.26	0.29	0.16	AGN2
CXOSEXSI J084905.3+445203	LynxE	1.266	24.61	43.8	1.11	1.27	0.74	ELG
CXOSEXSI J084831.6+445442	LynxW	1.267	25.42	43.2	0.61	1.23	0.8	ELG

NOTE. — AGN in high-redshift clusters of galaxies. Columns are: (1) AGN name; (2) Cluster; (3) AGN redshift; (4)  $R$ -band magnitude; (5) Rest-frame, hard-X-ray luminosity (2–10 keV); (6) Velocity offset from the cluster systemic velocity normalized by the cluster velocity dispersion; (7) Projected radial offset relative to the centroid of the X-ray gas in arcminutes; (8) Projected radial offset normalized by the cluster virial radius; (9) Spectroscopic classification. The  $R$ -band magnitude of E0015+162 is from Örndahl et al. (2003). The remaining values are from either Eckart et al. (2006) for the SESXI sample or from Silverman et al. (2005) for the ChaMP sample (although corrected from  $r'$  to  $R$  as noted in Section 2). The 2–8 keV X-ray luminosities from Silverman et al. (2005) have been corrected to the 2–10 keV band as described in Section 2.

We also correct the ChaMP  $r'$  measurements to the Vega  $R$  band as discussed above. Based on the magnitudes of these sources and a simple  $k$ -correction, we estimate that none of these sources falls below our galaxy luminosity threshold. As these are fairly luminous AGN, in some cases the AGN may dominate the total flux and we may have overestimated the host galaxy luminosity. E0015+162 (Margon et al. 1983) is the most X-ray luminous AGN in our sample by over an order of magnitude and is a useful case study to test the importance of this concern. This AGN has a total  $R = 18.41$  mag and a fainter host galaxy magnitude of  $R = 19.8$  mag (Örndahl et al. 2003), which corresponds to a factor of 3.6 in flux. If the other AGN have similar or smaller  $L_R/L_X$  ratios (such as due to obscuration), then we expect their AGN contribution to the measured  $R$ -band flux to be negligible because they are all much less luminous than E0015+162.

We identify AGN in these clusters with the following four criteria: 1) The hard X-ray luminosity must be  $L_{X,H} \geq 10^{43}$  erg s<sup>-1</sup>; 2) The AGN redshift must fall within  $3\sigma$  of the cluster mean redshift, where  $\sigma$  is the cluster velocity dispersion; 3) The AGN must fall within the projected virial radius  $R_{200}$  of the cluster; 4) The absolute magnitude of the host galaxy must be greater than  $M_R = M_R^*(z) + 1$  mag. Most of these criteria were adopted from Eastman et al. (2007), although the absolute magnitude criterion is different and we discuss our motivation for this choice in §4.3 below. With these criteria we identify 18 AGN in the 15 clusters with  $z > 0.4$ , or an average of more than one per cluster. The properties of the  $z > 0.4$  AGN are presented in Table 2.

TABLE 3  
NEW LOW-REDSHIFT CLUSTERS

Cluster (1)	$\alpha_c$ (2)	$\delta_c$ (3)	$z$ (4)	$\sigma$ [km/s] (5)	$\sigma$ Ref (6)	$T_X$ [keV] (7)	$T_X$ Ref (8)	$R_{200}$ [Mpc] (9)
Abell 1240	11:23:37.3	+43:06:54	0.1590	698	1	...	...	1.64
Abell 1942	14:38:22.0	+03:40:07	0.2240	903	2	5.6	1	1.96
Abell 2125	15:41:13.2	+66:16:01	0.2465	1113	3	3.2	2	2.39
MS1455.0+2232	14:57:15.1	+22:20:29	0.2578	1032	4	5.5	3	2.20
ZwCl 1358.1+6245	13:59:50.6	+62:31:04	0.3280	1003	4	6.5	3	2.06
MS1512.4+3647	15:14:22.4	+36:36:21	0.3720	575	4	3.6	3	1.15

NOTE. — New low-redshift clusters and their properties derived from the present study. Columns are: (1) Cluster name; (2 and 3) RA and DEC for the centroid of the extended X-ray emission; (4) redshift; (5) velocity dispersion; (6) reference for the velocity dispersion; (7) X-ray temperature in keV; (8) reference for the X-ray temperature; (9) estimate of the virial radius in Mpc (Treu et al. 2003). References for velocity dispersion are: 1: derived from the X-ray luminosity following Xue & Wu (2000); 2: derived from the X-ray temperature following Xue & Wu (2000); 3: Miller et al. (2004); 4: Borgani et al. (1999). References for X-ray temperatures are: 1: Ota & Mitsuda (2004); 2: Wang et al. (2004); 3: Mushotzky & Scharf (1997).

TABLE 4  
*Chandra* OBSERVATION LOGS

Cluster (1)	OBSID (2)	Detector (3)	T (ks) (4)	$L_{X,H, \text{Lim}}$ ( $10^{41} \text{ erg s}^{-1}$ ) (5)
Abell 1240	4961	ACIS-I	51.3	1.2
Abell 1942	3290	ACIS-I	57.5	2.2
Abell 2125	2207	ACIS-I	81.5	1.9
MS 1455.0+2232	4192	ACIS-I	91.9	1.8
ZwCl 1358.1+6245	516	ACIS-S3	53.0	2.6
MS 1512.4+3647	800	ACIS-S3	36.4	4.6

NOTE. — *Chandra* Observation Log. Columns are: (1) Cluster targeted; (2) Observation ID of *Chandra* data; (3) Detector used; (4) Usable exposure; (5) Estimate of the 2.0–8.0 keV luminosity limit of the observation for a cluster galaxy.

AGN more luminous than  $L_{X,H} = 10^{43} \text{ erg s}^{-1}$  are sufficiently rare in low-redshift clusters that Poisson uncertainties (as opposed to sources of systematic errors) from the low-redshift sample may dominate the statistical significance of any evidence of evolution. Our previous study of ten clusters with  $z < 0.32$  only identified one AGN above this luminosity threshold (Martini et al. 2006), while our more recent observations of three additional clusters (all at  $z < 0.08$ ) have identified only one additional AGN above this luminosity (Sivakoff et al. 2008). We have therefore studied six additional clusters with  $0.15 < z < 0.4$  to find other X-ray AGN more luminous than  $L_{X,H} = 10^{43} \text{ erg s}^{-1}$  with a combination of *Chandra* archival data and follow-up spectroscopy of candidate cluster X-ray AGN at the MDM Observatory. These clusters were selected to be the nearest massive clusters in the *Chandra* archive whose estimated virial radii fit within the *Chandra* ACIS field of view (FOV) and were accessible during our observing runs. The new clusters and their physical properties are listed in Table 3.

### 3.1. *Chandra* X-ray Analysis

The X-ray observations were processed following the same techniques employed by Sivakoff et al. (2008). We reduced all data using CIAO 3.4<sup>4</sup> with CALDB 3.3.0.1 and NASA’s FTOOLS 6.0<sup>5</sup>. The observations are summarized in Table 4. Only minor differences in reduction were required for these archival observations. The majority of the clusters had data with an aimpoint centered on the four ACIS-I chips ( $\sim 17'$  FOV) and frame times of 3.1 s. These data were telemetered and cleaned in Very Faint mode. The more distant clusters, ZwCl 1358.1+6245 and MS 1512.4+3647, were observed

with the aimpoint placed on the ACIS-S3 detector (8.4' FOV) and had frame times of 3.3 s. Their data were telemetered and cleaned in Faint mode, and thus have a slightly higher background. As all observations were operated at  $-120^\circ\text{C}$  the X-ray data were corrected for the time dependence of the gain and the charge-transfer inefficiency with their photon energies determined using the gain file acisD2000-01-29gain\_ctiN0006.fits. The archival data of all observations already had applied the newest tools to detect hot pixels and cosmic ray afterglows. We only consider events with ASCA grades of 0, 2, 3, 4, and 6. Known aspect offsets were applied for each observation. All observations were corrected for quantum efficiency degradation and had exposure maps determined at 1.5 keV. We excluded bad pixels, bad columns, and columns adjacent to bad columns or chip node boundaries. We also filtered out times when the blank-sky rate was more than three times the expected blank-sky rate derived from calibrated blank-sky backgrounds to avoid the most extreme periods of high background (“background flares”) that *Chandra* may encounter. MS 1512.4+3647 had two separate pointings and this introduced difficulties into our standard processing. We therefore excluded the shorter second pointing, which accounted for less than 25% of the total integration time.

To detect X-ray sources that are potential X-ray AGN in these clusters, we applied the wavelet detection algorithm (CIAO WAVDETECT) with scales ranging from 1 to 64 pixels in steps of  $\sqrt{2}$  factors and required a source detection threshold of  $10^{-6}$ . Source detection was only performed in regions with an exposure of greater than 10% of the total for the observation. Our source detection threshold corresponds to  $\lesssim 4$  falsely detected X-ray sources (due to a statistical fluctuation) for each observation. Using Kim et al. (2007), we have estimated the statistical X-ray positional uncertainty ( $1\sigma$ ) due to WAVDETECT. In Table 4, we list an estimated limiting X-ray luminosity for each observation that corresponds to five counts on axis (for consistency with Martini et al. 2006). For our analysis we concentrated on sources with at least 20 broad (0.3–8.0 keV) X-ray counts. These sources are unlikely to be due to statistical fluctuations except where they are coincident with ICM emission.

We used ACIS Extract 3.131<sup>6</sup> to create source extraction regions enclosing 90% of the flux in the X-ray PSF and to determine a masking radius that encircled 97% of the flux. For most of the sources, whose photons had median energies of  $\sim 0.6$ – $2.6$  keV, we determined the regions assuming the PSF at 1.497 keV. A few sources had harder emission and their PSF was calculated assuming an energy of 4.51 keV. In a

<sup>4</sup> <http://asc.harvard.edu/ciao/>.

<sup>5</sup> <http://heasarc.gsfc.nasa.gov/docs/software/lheasoft/>.

<sup>6</sup> [http://www.astro.psu.edu/xray/docs/TARA/ae\\_users\\_guide.html](http://www.astro.psu.edu/xray/docs/TARA/ae_users_guide.html)

relatively small number of crowded regions, the PSF fraction was reduced to prevent overlapping source extraction regions. We also used ACIS Extract to correct the (CIAO WAVDETECT) position to the mean position of detected events for sources within  $5'$  of the observation aimpoint or to the position that best correlated with the PSF for sources beyond  $5'$  of the observation aimpoint. These new positions were registered with an optical catalog from  $R$ -band images (see below) to correct the absolute astrometry and determine the absolute astrometric precision of each *Chandra* observation ( $0.3$ – $0.5''$ ). The statistical significance of each detection was added in quadrature with the absolute astrometric precision to estimate the total X-ray positional precision. We measured the counts in three energy ranges: the broad ( $0.3$ – $8$  keV), soft ( $0.3$ – $2$  keV), and hard ( $2.0$ – $8.0$  keV) bands. The observed fluxes in these bands were derived assuming a  $\Gamma = 1.7$  power-law spectrum with Galactic absorption. We then calculated the rest-frame luminosity in the broad band ( $0.3$ – $8$  keV) and the classic hard band ( $2$ – $10$  keV) for all sources with redshifts (see §3.3).

### 3.2. MDM Photometry

$R$ -band images of these clusters were obtained at the MDM Observatory 2.4m Hiltner telescope with the Echelle CCD camera during a run from the night of 28 May 2007 to 3 June 2007. Because the FOV of the CCD camera ( $\sim 9.5' \times 9.5'$ ) is smaller than the ACIS-I FOV ( $\sim 17' \times 17'$ ), we imaged a  $2 \times 2$  mosaic to cover the *Chandra* area, with each panel consisting of  $3 \times 300$ s exposures. All images were trimmed, bias-subtracted and flat-fielded with the CCDPROC package within IRAF<sup>7</sup>.

Sources were cataloged with the SExtractor package (Bertin & Arnouts 1996). Aperture magnitudes from these catalogs were calibrated with multiple observations of standard star fields from the data compiled by P. B. Stetson<sup>8</sup> onto the Vega magnitude system. Only data from the last night, which includes each quadrant of Abell 1240 and ZwCl 1358.1+6245, the north-east quadrant of MS 1512.4+3647, and  $1 \times 300$ s exposures of each quadrant of Abell 2125, were taken under photometric conditions. Our derived photometric solution for this night was precise to 0.03 mag. As all of these clusters except for Abell 2125 were imaged with SDSS, we cross-correlated aperture magnitudes from all images on this run with stars in the SDSS DR5 catalog. After correcting to  $R$  (Vega)<sup>9</sup>, our derived photometric solution for 3 June, which includes a color correction term, is accurate to 0.01 mag and precise to 0.06 mag. The poorer precision compared to our photometric solution appears to be only partially due to the dispersion in the Vega correction (overlapping sources between quadrants of our own observations indicate typical photometric precisions of  $0.05$ – $0.08$  mag). We therefore adopted the SDSS cross-calibration technique to photometrically correct all observations on non-photometric nights, except for observations of Abell 2125. For Abell 2125, non-photometric observations were cross calibrated with the single photometric exposures for Abell 2125. As we do not have complete multi-band data, we report only the magnitudes assuming no color correction.

<sup>7</sup> IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy (AURA) under cooperative agreement with the National Science Foundation.

<sup>8</sup> <http://cadwww.hia.nrc.ca/standards>

<sup>9</sup> <http://www.sdss.org/dr7/algorithms/sdssUBVRITransform.html#Lupton2005>

The exclusion of the color correction term does not significantly decrease the precision of our photometric solutions.

We calculated astrometric solutions for the images with the WCSTools package (Mink 2002), package and then produced the final, calibrated mosaics with the SWARP<sup>10</sup> package. A final source catalog was extracted with SExtractor and used to register the astrometry of the X-ray observations. We consider only the SExtractor AUTO magnitudes, which is an automatic aperture magnitude designed to give precise estimates of total magnitudes for galaxies. As nearby, detected neighbors are removed and replaced by mirroring the opposite side of the aperture where available, these magnitudes are suitable for our relatively crowded fields. All X-ray sources that would be more luminous than  $L_{X,H} = 10^{43}$  erg s<sup>-1</sup> at the cluster redshift that were also associated with galaxies and that would be more luminous than  $M_R^*(z) + 1$  at the cluster redshift were then targeted for the highest-priority spectroscopic observations, with the exception of sources heavily contaminated by ICM emission. We also identified other candidate cluster X-ray AGN, specifically those that would have  $L_{X,H} \geq 10^{42}$  erg s<sup>-1</sup>, as lower-priority spectroscopic targets.

### 3.3. MDM Spectroscopy

We obtained low-resolution spectroscopy of these candidates with the 2.4m Hiltner telescope with the CCDs, a Boller & Chivens spectrograph, during a run from the night of 28 April 2008 to 3 May 2008. The slit widths were determined by the nightly seeing conditions and were either  $1.0''$  or  $1.5''$ . At least two exposures of every candidate were obtained and total exposure times varied from 120s to 9000s. Five sets of internal and twilight flats were taken over the entire run, while comparison lamps were observed before and/or after every candidate.

The files were trimmed and bias-subtracted with the CCDPROC package within IRAF and bad pixels were determined from a ratio of flat-field images and were fixed in every image. The individual flat-field images from internal lamps revealed a complex wavelength and slit-dependent flat-field, most likely due to some reflection. To model this complex response, we first median smoothed the internal flat-fields (over  $11 \times 11$  pixels) and then Gaussian-smoothed ( $\sigma = 11$  pixels) over the dispersion axis. The ratio of the internal flat-field to the modeled internal flat-field was adopted as the true internal flat-field. An illumination correction was then created from the twilight flat-fields and applied to make the final set of flat-field corrections to remove fringing in the spectra. After each spectrum was properly flat-fielded, we rejected cosmic rays using L.A. Cosmic<sup>11</sup> (van Dokkum 2001). A fourth order wavelength-solution was calculated for each set of HgNe comparison spectra, resulting in a typical RMS of  $\sim 0.1 \text{ \AA pixel}^{-1}$ . Thereafter, standard aperture extraction of the spectra was used to remove the night sky emission and produce one-dimensional, logarithmically interpolated spectra with a dispersion of  $\sim 3 \text{ \AA pixel}^{-1}$ . The spectra extend from approximately  $3650 \text{ \AA}$  to  $7250 \text{ \AA}$ . We extracted both the signal and noise for each final spectrum of a source.

We adapted the Princeton/MIT SDSS Spectroscopy routines<sup>12</sup> to calculate redshifts. This technique cross-correlates the spectra in pixel space with template spectra, with each

<sup>10</sup> [http://terapix.iap.fr/rubrique.php?id\\_rubrique=49](http://terapix.iap.fr/rubrique.php?id_rubrique=49)

<sup>11</sup> <http://www.astro.yale.edu/dokkum/lacosmic/>

<sup>12</sup> [http://spectro.princeton.edu/idspec2d\\_doc.html](http://spectro.princeton.edu/idspec2d_doc.html)

TABLE 5  
NEW, LOWER-LUMINOSITY CLUSTER X-RAY AGN

CXOU ID (1)	$z$ (2)	$z$ ref (3)	$R$ (4)	$R$ flag (5)	$f_{X,S}$ (6)	$f_{X,H}$ (7)	$f_{X,B}$ (8)	$L_{X,B}$ (9)	$L_{X,H}$ (10)	X flag (11)
J135950.5+623106.3	$0.32717 \pm 0.00038$	1	$17.80 \pm 0.05$	3	$9.5^{+1.1}_{-1.0}$	$8.7^{+2.8}_{-2.4}$	$20.6^{+2.2}_{-2.1}$	$70.6^{+7.6}_{-7.2}$	$46.2^{+5.0}_{-4.7}$	1
J143821.8+034013.3	0.22479	2	$16.44 \pm 0.06$	3	$2.20^{+0.91}_{-0.76}$	$2.3^{+1.9}_{-1.4}$	$4.8^{+1.7}_{-1.5}$	$7.2^{+2.5}_{-2.2}$	$4.7^{+1.7}_{-1.4}$	1
J145714.7+221933.6	$0.24852 \pm 0.00025$	1	$20.04 \pm 0.07$	0	$2.40^{+0.63}_{-0.53}$	$3.7^{+1.3}_{-1.1}$	$5.9^{+1.2}_{-1.0}$	$11.2^{+2.2}_{-2.0}$	$7.3^{+1.5}_{-1.3}$	0
J145715.0+222034.5	$0.25772 \pm 0.00015$	1	$16.82 \pm 0.07$	0	$20.2^{+4.9}_{-4.8}$	$21.6^{+8.5}_{-8.3}$	$44.4^{+9.1}_{-9.0}$	$93^{+19}_{-19}$	$61^{+12}_{-12}$	1
J151422.5+363620.7	0.3718	3	$18.05 \pm 0.06$	2	$3.98^{+0.98}_{-0.89}$	$3.2^{+2.3}_{-1.9}$	$8.4^{+1.8}_{-1.8}$	$38.1^{+8.8}_{-8.0}$	$24.9^{+5.7}_{-5.2}$	1
J154101.9+661627.1	$0.24564 \pm 0.00045$	1	$17.19 \pm 0.08$	2	$2.78^{+0.62}_{-0.57}$	$0.21^{+0.76}_{-0.41}$	$4.63^{+1.0}_{-0.87}$	$8.5^{+1.9}_{-1.6}$	$5.5^{+1.2}_{-1.0}$	0
J154101.9+661721.4	0.2567	4	$19.36 \pm 0.08$	0	$8.11^{+0.97}_{-0.88}$	$7.1^{+1.7}_{-1.4}$	$17.0^{+1.8}_{-1.6}$	$34.3^{+3.6}_{-3.3}$	$22.4^{+2.3}_{-2.1}$	0
J154117.3+661923.6	0.2465	4	$18.81 \pm 0.08$	0	$2.08^{+0.38}_{-0.47}$	$1.46^{+1.2}_{-0.82}$	$4.15^{+1.1}_{-0.88}$	$7.6^{+1.9}_{-1.6}$	$5.0^{+1.3}_{-1.1}$	0

NOTE. — *Chandra* Observation Log. Columns are: (1) Name of X-ray source; (2) Redshift (3) References for redshift are: 1: this work; 2: SDSS (Adelman-McCarthy 2008); 3 Abraham et al. (1998); 4: Miller et al. (2004); (4)  $R$ -band magnitude; (5) Flags for photometry are: (0) no flag; (1) may be contaminated by nearby neighbors or bad pixels; (2) blended with nearby neighbors; (3) both; (6–8) Soft [0.3–2 keV], Hard [2–8 keV], and Broad [0.3–8 keV] band flux in the observed frame in units of  $10^{-15} \text{ erg s}^{-1} \text{ cm}^{-2}$ . (9–10) Broad [0.3–8 keV] and Hard [2–10 keV] band luminosity in the rest-frame in units of  $10^{41} \text{ erg s}^{-1}$  corrected for Galactic absorption. (11) X-ray flags are: (0) no flag; (1) contaminated by ICM peak. Note that CXOU J145715.0+222034.5 is the BCG and we subtracted a multi-component beta model for the ICM to compute the quoted fluxes and luminosities.

pixel weighted by the inverse of its variance, and is similar to the technique used in Martini et al. (2006). The template spectra include a set of four eigenspectra for galaxies, four eigenspectra for quasars, and forty eigenspectra for stars. The five best galaxy redshifts for  $-0.01 < z < 1.00$ , five best quasar redshifts for  $0.0033 < z < 7.00$ , and forty different stellar redshifts for  $-0.004 < z < 0.004$  are found and ordered by the reduced  $\chi^2$  of their fit. We adopted the best-fit redshift and classification for each source. To ascertain the quality of the fit and the errors to the redshift, we resampled each spectra 100 times randomly according to its noise characteristics and reran the cross-correlation routine. Both the dispersion in best-fit redshifts and the best-fit spectral type were used to qualify the spectral classification quality. If the dispersion in redshift was relatively low ( $\sigma_z \lesssim 0.01$ ),  $> 68\%$  of the best-fit redshifts were within  $3\sigma_z$  of our adopted redshift, and had the same spectral type (i.e., galaxy, quasar, or a similar stellar type) we consider this a secure redshift. Typically the maximum SNR of these spectra were  $> 5 \text{ pixel}^{-1}$ .

We did not identify any AGN in these clusters with  $L_{X,H} \geq 10^{43} \text{ erg s}^{-1}$ , although we did identify several lower-luminosity AGN in these clusters. Data for the lower-luminosity X-ray sources are provided in Table 5 and include several sources with spectroscopic measurements from the literature. The spectroscopic observations of Abell 1240 and MS1512.4+3647 are complete for all candidates that would have  $L_{X,H} \geq 10^{42}$  if at the cluster redshift, while the other four clusters are not complete to this luminosity limit. We have also measured redshifts,  $R$ -band magnitudes, and X-ray fluxes and luminosities for numerous additional sources not associated with these clusters and their properties are listed in Table 6. As for the high-redshift clusters, several of the low-redshift clusters do not have direct velocity dispersion measurements. For Abell 1942 we estimated this quantity from the X-ray temperature. For Abell 1240 Xue & Wu (2000) quote  $kT = 3.83 \text{ keV}$  from (Mushotzky & Scharf 1997), but in fact the value in Mushotzky & Scharf (1997) appears instead to be for Abell 1242. As we could not identify another  $T_X$  value in the literature, we used the measurement of  $L_{\text{bol}} = 2.71 \times 10^{44} \text{ erg s}^{-1}$  from David et al. (1999) and the relation  $\sigma = 10^{2.76} L_X^{0.19}$  derived by Xue & Wu (2000) to estimate the velocity dispersion.

#### 4. CLUSTER X-RAY AGN FRACTION

We require two quantities to estimate the AGN fraction in these clusters: the number of AGN above our hard X-ray luminosity threshold hosted by galaxies with  $M_R < M_R^*(z) + 1$  and the total number of cluster galaxies above this magnitude threshold. For our low-redshift cluster sample, we have complete data to our X-ray threshold and reasonably complete data for the other cluster galaxies for about half of the clusters. For the high-redshift sample we have incomplete knowledge of both quantities. The AGN sample is likely incomplete because of spectroscopic incompleteness in the ChaMP and SEXSI surveys. The census of other cluster galaxies is very incomplete because few very high redshift clusters have the same quality membership data as our low-redshift sample. In the first three subsections below we describe the choice of the fiducial absolute magnitude threshold, our estimate of the completeness of the spectroscopic observations of X-ray sources, and the total number of cluster galaxies in the clusters with incomplete membership data. The fourth subsection describes our main result, the measurement of the AGN fraction and its evolution. The final two subsections describe potential contamination by AGN associated with large-scale structure around these clusters and other sources of uncertainty, respectively.

##### 4.1. Host galaxy magnitude threshold

In previous work we defined the AGN fraction in clusters relative to galaxies more luminous than an  $R$ -band absolute magnitude of  $M_R = -20 \text{ mag}$  (e.g. Martini et al. 2006). This choice of magnitude threshold was largely driven by expedience, namely it corresponded to the completeness limit for the most distant clusters in that sample. To properly extend this work to high redshift it is important to account for the evolution of the galaxy population in clusters, both in luminosity and number. These were not significant effects in our low-redshift study as the highest-redshift cluster was at only  $z = 0.31$ , but in our previous work at  $z \sim 0.6$  by Eastman et al. (2007) the  $M_R = -20 \text{ mag}$  cutoff corresponded to a fainter absolute magnitude relative to  $M_R^*$ . Because the cluster galaxy population is larger, this would have led to a lower estimate of the AGN fraction if the cluster AGN are predominantly associated with the most luminous galaxies, as is the case at low redshifts (Sivakoff et al. 2008).

Here we adopt an absolute magnitude threshold of  $M_R^*(z) + 1$ , and thus allow for evolution of  $M_R^*$ . At low-redshifts ( $0.01 < z < 0.07$ ) Christlein & Zabludoff (2003) measured

TABLE 6  
 NONMEMBER X-RAY SOURCES

CXOU ID (1)	$z$ (2)	$z$ ref (3)	$R$ (4)	$R$ flag (5)	$f_{X,S}$ (6)	$f_{X,H8}$ (7)	$f_{X,B}$ (8)	$\log L_{X,B}$ (9)	$\log L_{X,H}$ (10)
J112314.9+431208.3	0.08017 $\pm$ 0.00010	1	17.66 $\pm$ 0.08	0	8.4 <sup>+1.5</sup> <sub>-1.3</sub>	29.9 <sup>+4.7</sup> <sub>-4.1</sub>	30.1 <sup>+3.4</sup> <sub>-3.1</sub>	41.69 <sup>+0.05</sup> <sub>-0.04</sub>	41.50 <sup>+0.05</sup> <sub>-0.03</sub>
J112357.4+431314.1	0.08007	2	19.46 $\pm$ 0.08	0	23.8 <sup>+2.2</sup> <sub>-2.2</sub>	32.8 <sup>+4.9</sup> <sub>-4.3</sub>	55.9 <sup>+3.5</sup> <sub>-3.2</sub>	44.51 <sup>+0.03</sup> <sub>-0.03</sub>	44.32 <sup>+0.03</sup> <sub>-0.03</sub>
J112403.0+431330.6	1.1049	2	18.39 $\pm$ 0.08	0	22.2 <sup>+2.5</sup> <sub>-2.2</sub>	17.3 <sup>+4.2</sup> <sub>-3.6</sub>	44.5 <sup>+4.4</sup> <sub>-4.0</sub>	43.16 <sup>+0.04</sup> <sub>-0.04</sub>	42.98 <sup>+0.04</sup> <sub>-0.04</sub>
J112413.1+430639.3	2.3666 $\pm$ 0.0015	1	19.80 $\pm$ 0.08	0	7.53 <sup>+1.3</sup> <sub>-1.1</sub>	7.56 <sup>+2.5</sup> <sub>-2.0</sub>	16.1 <sup>+2.4</sup> <sub>-2.1</sub>	44.73 <sup>+0.07</sup> <sub>-0.06</sub>	44.54 <sup>+0.07</sup> <sub>-0.06</sub>
J143804.9+033752.6	0.29192 $\pm$ 0.00030	1	18.50 $\pm$ 0.06	0	5.22 <sup>+1.1</sup> <sub>-0.95</sub>	< 4.0	7.8 <sup>+1.8</sup> <sub>-1.5</sub>	42.32 <sup>+0.10</sup> <sub>-0.09</sub>	42.13 <sup>+0.10</sup> <sub>-0.09</sub>
J143832.2+033506.0	1.0083 $\pm$ 0.0051	1	19.98 $\pm$ 0.06	0	65.7 <sup>+6.2</sup> <sub>-3.0</sub>	79.8 <sup>+6.0</sup> <sub>-3.8</sub>	149.3 <sup>+6.0</sup> <sub>-3.8</sub>	44.85 <sup>+0.02</sup> <sub>-0.02</sub>	44.66 <sup>+0.02</sup> <sub>-0.02</sub>
J143833.0+033606.8	0.38252 $\pm$ 0.00017	1	19.40 $\pm$ 0.06	0	11.5 <sup>+1.2</sup> <sub>-1.2</sub>	18.3 <sup>+2.7</sup> <sub>-2.7</sub>	28.5 <sup>+2.8</sup> <sub>-2.4</sub>	43.15 <sup>+0.04</sup> <sub>-0.04</sub>	42.96 <sup>+0.04</sup> <sub>-0.04</sub>
J143839.7+033631.3	2.1493 $\pm$ 0.0019	1	19.00 $\pm$ 0.06	0	16.1 <sup>+1.7</sup> <sub>-1.5</sub>	16.1 <sup>+3.1</sup> <sub>-2.6</sub>	34.8 <sup>+3.1</sup> <sub>-2.8</sub>	44.97 <sup>+0.04</sup> <sub>-0.04</sub>	44.79 <sup>+0.04</sup> <sub>-0.04</sub>
J143841.9+034110.2	1.7372	2	17.82 $\pm$ 0.06	0	28.2 <sup>+2.3</sup> <sub>-2.2</sub>	29.5 <sup>+4.3</sup> <sub>-3.8</sub>	61.6 <sup>+4.3</sup> <sub>-3.8</sub>	45.01 <sup>+0.03</sup> <sub>-0.03</sub>	44.83 <sup>+0.03</sup> <sub>-0.03</sub>
J143847.3+032950.8	0.00034 $\pm$ 0.00012	1	16.89 $\pm$ 0.06	0	16.3 <sup>+2.0</sup> <sub>-1.8</sub>	2.0 <sup>+3.8</sup> <sub>-1.6</sub>	26.9 <sup>+3.0</sup> <sub>-3.0</sub>		
J143859.0+033547.8	0.7339	2	18.51 $\pm$ 0.06	0	46.5 <sup>+2.8</sup> <sub>-2.7</sub>	7.1 <sup>+6.1</sup> <sub>-5.7</sub>	113.8 <sup>+5.6</sup> <sub>-5.3</sub>	44.41 <sup>+0.02</sup> <sub>-0.02</sub>	44.22 <sup>+0.02</sup> <sub>-0.02</sub>
J145623.0+221833.5	0.00027 $\pm$ 0.00010	1	15.51 $\pm$ 0.07	0	9.0 <sup>+1.5</sup> <sub>-1.3</sub>	5.1 <sup>+2.5</sup> <sub>-2.0</sub>	17.1 <sup>+2.5</sup> <sub>-2.0</sub>		
J145624.5+222057.1	0.00019 $\pm$ 0.00010	1	15.45 $\pm$ 0.07	0	14.8 <sup>+1.4</sup> <sub>-1.3</sub>	11.3 <sup>+2.3</sup> <sub>-2.2</sub>	29.8 <sup>+2.3</sup> <sub>-2.2</sub>		
J145634.6+221514.2	0.40918 $\pm$ 0.00010	1	20.16 $\pm$ 0.07	0	25.8 <sup>+1.7</sup> <sub>-1.6</sub>	56.8 <sup>+4.4</sup> <sub>-4.1</sub>	73.3 <sup>+3.7</sup> <sub>-3.3</sub>	43.63 <sup>+0.02</sup> <sub>-0.02</sub>	43.45 <sup>+0.02</sup> <sub>-0.02</sub>
J145657.7+221315.6	0.00016 $\pm$ 0.00010	1	14.87 $\pm$ 0.07	0	8.63 <sup>+1.0</sup> <sub>-0.93</sub>	3.6 <sup>+1.4</sup> <sub>-1.1</sub>	16.0 <sup>+1.6</sup> <sub>-1.6</sub>		
J145708.7+222352.4	0.1238	2	17.44 $\pm$ 0.07	0	2.27 <sup>+0.60</sup> <sub>-0.50</sub>	< 3.4	3.32 <sup>+1.0</sup> <sub>-0.86</sub>	41.14 <sup>+0.13</sup> <sub>-0.11</sub>	40.95 <sup>+0.13</sup> <sub>-0.11</sub>
J145710.7+221844.9	1.885 $\pm$ 0.0014	1	18.73 $\pm$ 0.07	0	3.99 <sup>+0.66</sup> <sub>-0.57</sub>	5.3 <sup>+1.4</sup> <sub>-1.1</sub>	9.4 <sup>+1.2</sup> <sub>-1.1</sub>	44.28 <sup>+0.06</sup> <sub>-0.05</sub>	44.09 <sup>+0.06</sup> <sub>-0.05</sub>
J145712.3+221446.7	-0.00069 $\pm$ 0.00010	1	15.15 $\pm$ 0.07	1	50.4 <sup>+2.1</sup> <sub>-2.1</sub>	15.3 <sup>+2.1</sup> <sub>-2.0</sub>	90.2 <sup>+3.6</sup> <sub>-3.5</sub>		
J145721.0+222334.5	1.7362 $\pm$ 0.0010	1	19.33 $\pm$ 0.07	1	9.4 <sup>+1.1</sup> <sub>-1.0</sub>	7.0 <sup>+1.9</sup> <sub>-1.6</sub>	18.9 <sup>+2.0</sup> <sub>-1.9</sub>	44.50 <sup>+0.05</sup> <sub>-0.04</sub>	44.32 <sup>+0.05</sup> <sub>-0.04</sub>
J145726.9+221755.1	1.4664 $\pm$ 0.0011	1	19.55 $\pm$ 0.07	0	23.6 <sup>+1.7</sup> <sub>-1.6</sub>	33.0 <sup>+3.4</sup> <sub>-3.1</sub>	56.3 <sup>+3.2</sup> <sub>-3.0</sub>	44.81 <sup>+0.03</sup> <sub>-0.02</sub>	44.62 <sup>+0.03</sup> <sub>-0.02</sub>
J151427.0+363803.1	0.1616	2	16.90 $\pm$ 0.06	0	2.28 <sup>+0.61</sup> <sub>-0.49</sub>	1.8 <sup>+1.9</sup> <sub>-1.1</sub>	4.82 <sup>+1.2</sup> <sub>-0.99</sub>	41.53 <sup>+0.11</sup> <sub>-0.09</sub>	41.35 <sup>+0.11</sup> <sub>-0.09</sub>
J151428.4+363743.5	0.4026	3	20.13 $\pm$ 0.06	0	7.70 <sup>+1.0</sup> <sub>-0.92</sub>	14.1 <sup>+3.7</sup> <sub>-3.0</sub>	18.9 <sup>+2.2</sup> <sub>-2.0</sub>	43.01 <sup>+0.05</sup> <sub>-0.05</sub>	42.83 <sup>+0.05</sup> <sub>-0.05</sub>
J151437.5+364041.3	0.1468	3	19.86 $\pm$ 0.06	0	12.1 <sup>+1.2</sup> <sub>-1.1</sub>	13.5 <sup>+3.5</sup> <sub>-2.9</sub>	26.9 <sup>+2.4</sup> <sub>-2.3</sub>	42.19 <sup>+0.04</sup> <sub>-0.04</sub>	42.01 <sup>+0.04</sup> <sub>-0.04</sub>
J153938.1+662102.4	0.4375	4	19.71 $\pm$ 0.08	0	5.02 <sup>+1.2</sup> <sub>-0.98</sub>	6.5 <sup>+2.9</sup> <sub>-2.4</sub>	11.7 <sup>+2.3</sup> <sub>-2.1</sub>	42.90 <sup>+0.09</sup> <sub>-0.08</sub>	42.71 <sup>+0.09</sup> <sub>-0.08</sub>
J154012.3+661439.2	1.0577 $\pm$ 0.0029	1	19.75 $\pm$ 0.08	0	37.2 <sup>+1.8</sup> <sub>-1.9</sub>	41.9 <sup>+3.7</sup> <sub>-3.4</sub>	83.1 <sup>+3.7</sup> <sub>-3.5</sub>	44.64 <sup>+0.02</sup> <sub>-0.02</sub>	44.46 <sup>+0.02</sup> <sub>-0.02</sub>

NOTE. — *Chandra* Observation Log. Columns are: Col (1) Name of X-ray source; Col (2) Redshift Col (3) References for redshift are: 1: this work; 2: SDSS (Adelman-McCarthy 2008); 3: Abraham et al. (1998); 4: Miller et al. (2004); Col (4)  $R$ -band magnitude; Col (5) Flags for photometry are: (0) no flag; (1) may be contaminated by nearby neighbors or bad pixels; Cols (6–8) Soft [0.5–2 keV], Hard [2–8 keV], and Broad [0.5–8 keV] band flux in the observed frame in units of  $10^{-15} \text{ erg s}^{-1} \text{ cm}^{-2}$ . Upper limits are  $3\sigma$  limits. Cols (9–10) Log of the Broad [0.5–8 keV] and Hard [2–10 keV] band luminosity in the rest-frame in units of  $\text{erg s}^{-1}$  corrected for Galactic absorption. We do not quote luminosities for X-ray sources identified with Galactic stars ( $z \sim 0$ ).

the  $R$ -band luminosity function (LF) for six nearby clusters<sup>13</sup> and found that the composite cluster LF is consistent with a Schechter function with  $M_R^* = -21.92 \pm 0.17 \text{ mag}$  ( $h = 0.7$ ,  $\alpha = -1.21$ ). They also find an essentially identical value of  $M_R^* = -21.93 \text{ mag}$  for the field. The low-redshift value of  $M_R^* + 1$  is therefore about one magnitude brighter than the value of  $M_R = -20 \text{ mag}$  we adopted in our previous, low-redshift studies (Martini et al. 2006; Sivakoff et al. 2008). For comparison, Blanton et al. (2003) measured  $M^* = -21.22$  ( $\alpha = -1.05$ ) at  $z = 0.1$  for the  $r^{0.1}$  band on the AB system. This corresponds to  $M_R^* = -21.72 \text{ mag}$  on the Vega system for the  $R$ -band at  $z = 0$  based on the conversions presented in Blanton & Roweis (2007) and is therefore consistent with Christlein & Zabludoff (2003).

Many recent studies have measured the evolution of  $M_R^*$  and generally these measurements include both a value for all galaxies and separate measurements for particular spectroscopic types. This has relevance for our study as the cluster galaxy population is on average more quiescent than field galaxies and consequently their evolutionary history is different. We are most interested in measurements of the evolution of  $M_R^*$  as a function of spectral type to isolate the evolution of galaxies dominated by older stellar populations that are most likely representative of the evolution of cluster galaxies. A useful, low-redshift benchmark for a type-dependent LF for clusters comes again from Christlein & Zabludoff (2003). They find  $M_R^* = -21.78 \text{ mag}$  for quiescent galaxies in clusters, which is nearly identical to the value for all cluster mem-

bers. For field galaxies Chen et al. (2003) use photometric redshifts in the Las Campanas Infrared Survey and measure values of  $-21.70$  to  $-22.22 \text{ mag}$  ( $\alpha = -1$ ) for all galaxies over the range  $0.5 < z < 1.5$  and values of  $-21.21$  to  $-21.82 \text{ mag}$  ( $\alpha = -0.2$ ) for galaxies consistent with an E/S0 + Sab spectral template. Wolf et al. (2003) use photometric redshifts from COMBO-17 and measure more pronounced evolution for their early-type spectral template with  $M_R^*$  fading by  $\sim 1 \text{ mag}$  from  $z \sim 1.1$  to  $z \sim 0.3$ . More recently, Ilbert et al. (2005) measure a fading of  $1.1$ – $1.8 \text{ mag}$  between  $z \sim 2$  and  $z \sim 0.1$  in the  $R$ -band based on spectroscopic redshifts, although they do not present the evolution as a function of spectral type. These measurements of evolution in  $M_R^*$  are broadly comparable to the  $1.2 \text{ mag}$  of fading from  $z = 1$  to the present expected from pure luminosity evolution of a single stellar population with  $z_f = 2$  and solar metallicity (Bruzual & Charlot 2003).

Direct measurements of evolution of the cluster LF have mostly been conducted in the rest-frame  $B$ -band. Goto et al. (2005) find  $M_B^* = -21.13 \text{ mag}$  for MS1054-03 ( $z = 0.83$ ), which is in our sample, and similar to the  $M_B^* = -21.15 \text{ mag}$  measured for three clusters at an average  $z = 0.859$  by Postman et al. (2001). In comparison to local  $B$ -band measurements of the cluster LF (e.g. Colless 1989; Rauzy et al. 1998), Goto et al. (2005) conclude that  $M_B^*$  fades by  $0.46$  to  $0.71 \text{ mag}$  between  $z = 0.83$  and  $z = 0$ . For the same simple stellar population model considered above (Bruzual & Charlot 2003),  $1.2 \text{ mag}$  of fading in  $B$ -band is expected from  $z = 0.83$  to the present. While there is not a direct measurement in the rest-frame  $R$ -band for the cluster LF, at yet longer wavelengths Ellis & Jones (2004) find that the fading in the

<sup>13</sup> Two of these clusters (Abell 85 and Abell 754) are in our low-redshift sample (Sivakoff et al. 2008).



$K$ -band is 1.2 mag from  $z = 0.9$  to the present and consistent with passive evolution and a formation epoch at  $z_f = 2$ . From these investigations of the LF evolution in the field and clusters, we adopt the assumption that  $M_R^*(z) = M_R^*(0) - z$  and the normalization for  $M_R^*$  from Christlein & Zabludoff (2003) for all cluster galaxies to estimate the completeness of the spectroscopy of X-ray counterparts and the size of the galaxy population in low-redshift clusters. This result is broadly consistent with all of the results described here, although is most consistent with the studies that predict more fading. If there is less fading of galaxies at the bright end of the LF, such as may be due to some low-level star formation in these galaxies, then the completeness limits we describe next are too bright and we will have systematically underestimated the population of luminous AGN in the higher-redshift clusters.

#### 4.2. Completeness

We calculate a completeness limit in the observed  $R$ -band for each cluster based on the value of  $M_R^*(z) + 1$  and a  $k$ -correction derived from the elliptical template of the four-component spectral template presented by Assef et al. (2008). These templates are derived from 16,033 galaxies with spectroscopic redshifts and multiband photometry from the AGN and Galaxy Evolution Survey. Most of the galaxies are in the range  $0 < z < 1$  and the median redshift is 0.31. The parent sample is therefore broadly representative of our redshift range. For the higher-redshift clusters the  $k$ -correction requires a substantial extrapolation from the observed  $R$ -band, which for example samples rest-frame  $B$ -band at  $z = 0.5$ . Our assumption that the typical cluster galaxies are best approximated by an elliptical template is certainly reasonable for the low-redshift clusters. This may not be as good an approximation at higher redshifts, although in a study of the color-magnitude relation in our two highest-redshift clusters (Lynx E and W) Mei et al. (2009) found there is no evidence for significant evolution. If a later-type template were a better choice for the  $k$ -correction at higher redshift, the  $k$ -correction would be smaller and the necessary  $R$ -band spectroscopic limit would be brighter. The net effect would be a smaller completeness correction.

The spectroscopic completeness of the high- $z$  AGN sample largely depends on the completeness of the ChaMP and SEXSI surveys, although we also use additional spectra for MS 2053.7-0449, MS 1054-03, and RDCS J0910+5422. The ChaMP survey quotes a spectroscopic completeness of 77% for  $R < 22.37$  mag (Silverman et al. 2005) and the SEXSI survey quotes a spectroscopic completeness of 61% for sources with  $22 < R < 23$  mag, 67% for sources with  $23 < R < 24$ , and 74% for sources with  $R > 24$  mag (typically to 24.4 mag) (Eckart et al. 2006). For the ChaMP data we adopt 77% as the completeness correction for  $R < 22.37$  mag, while for the SEXSI survey we adopt an average completeness correction of 67% for  $R < 24.4$  mag. For nearly all of the clusters above  $z > 0.6$  the spectroscopic data do not extend to the equivalent of  $M_R^*(z) + 1$  and the size of the magnitude range without spectra ranges from a few tenths to over a magnitude. To estimate the number that may have been missed we inspected the host galaxy absolute magnitude distribution of the  $L_{X,H} \geq 10^{43}$  erg s $^{-1}$  AGN in the clusters with complete data and find only one AGN fainter than  $M_R^*$ . The distribution in  $M_R$  of the X-ray AGN is shown in Figure 1. We therefore assume that we have not missed any AGN because the spectroscopic observations of X-ray sources did not have the requisite depth, although this assumption may have led us to underestimate the AGN

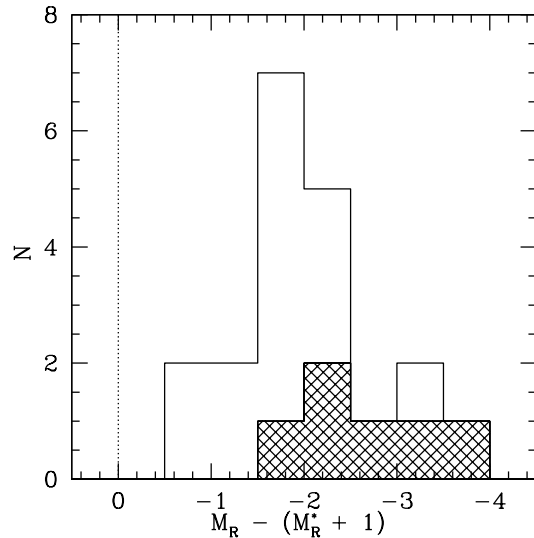


FIG. 1.— Distribution in absolute magnitude  $M_R$  of the cluster AGN relative to  $M_R^*(z) + 1$  at their redshift. All of the cluster AGN are substantially brighter than  $M_R^*(z) + 1$ , although in most cases the spectroscopy is complete to this limit. The subset that are classified as BLGN are represented by the hatched histogram. The dotted line corresponds to our galaxy luminosity threshold at  $M_R^*(z) + 1$ .

fraction at high redshift. In contrast, if our assumption of an early-type template for the  $k$ -correction was too red, then the spectroscopic data do achieve the requisite depth and this remains a nonissue. At brighter apparent magnitudes we do apply a completeness correction to account for the quoted 77% and 67% completeness of the surveys. We discuss this further in §4.4 below.

The X-ray AGN populations of several of these clusters have been studied in previous work. The first substantial study of spectroscopically-confirmed X-ray AGN in a high-redshift cluster was by Johnson et al. (2003) in MS1054-03. They identified 2 AGN associated with this cluster: CXOU J105702.7-033943 and CXOU J105710.6-033500; however, neither of these are included in the present sample. The first was not included because the X-ray luminosity is below our threshold of  $10^{43}$  erg s $^{-1}$  and the second because it falls slightly outside the projected virial radius ( $R/R_{200} = 1.2$ ). Martel et al. (2007) have also studied X-ray sources in clusters, including three clusters that overlap this sample. They are discussed further in §5.3 below.

#### 4.3. Inactive Cluster Galaxy Population

To estimate the AGN fraction in these clusters we need to know the number of cluster galaxies more luminous than  $M_R^*(z) + 1$ . We estimate this quantity in two ways, depending on the available data for the clusters. For the low-redshift clusters in our previous studies (Martini et al. 2006, 2007; Sivakoff et al. 2008) we have a large number of spectroscopically-confirmed cluster members and can estimate the number of cluster galaxies either directly or with a completeness correction. We have calculated new estimates for these clusters for the present paper because we no longer use the  $M_R = -20$  mag threshold of the previous work. These values are listed in Table 7.

For essentially all of the new clusters in the present study we employ the same technique as Eastman et al. (2007) to estimate the number of cluster members above  $M_R^*(z) + 1$  from the cluster velocity dispersion. This employs the richness-velocity dispersion relationship defined by Koester et al.

(2007) for the MaxBCG cluster sample. The cluster richness  $N_{gal}^{R200}$  is the number of red (E/S0) cluster members more luminous than  $0.4L^*$  within the projected  $R_{200}$  radius. This relationship was originally derived from a sample of 13,823 clusters with  $0.1 < z < 0.3$  in the SDSS with velocity dispersions greater than  $\sim 400 \text{ km s}^{-1}$ . Becker et al. (2007) provide the most recent estimate of this relation based on a larger sample that extends over both a broader redshift range and to lower velocity dispersion groups. They find  $\ln \sigma = (6.17 \pm 0.04) + (0.436 \pm 0.015) \ln N_{gal}^{R200}/25$ . For reference a  $520 \text{ km s}^{-1}$  cluster has  $N_{gal}^{R200} = 30$ .

There are several caveats that need to be considered with the use of this estimator. First, the richness–velocity dispersion relationship is based on photometric and not spectroscopic redshifts. This is not a significant concern because for red cluster galaxies the photometric redshift estimates are robust within the quoted uncertainties. The second concern is that this relationship is based on the red cluster galaxies alone. At low redshifts this estimate is a reasonable approximation as the vast majority of cluster galaxies more luminous than  $M_R^* + 1$  fall in this category. For example, the fraction of quiescent galaxies above this luminosity in the composite LF of Christlein & Zabludoff (2003) is  $\sim 85\%$ . While their definition of quiescence is based on spectral lines rather than color, these two definitions of quiescence typically agree when averaged over a cluster. At higher redshifts a larger fraction of the cluster galaxies may be blue due to ongoing star formation, but this can not be a substantial contribution because the luminosity-weighted mean star formation epoch is  $z = 2$  for early-type cluster galaxies up to  $z = 0.5$  (van Dokkum & van der Marel 2007). Becker et al. (2007) do find evidence of evolution in this relationship in the sense of lower richness at fixed velocity dispersion in higher redshift clusters, but they note that this may be due to their strict color selection. In addition, for our accounting of the inactive galaxy population the color of the galaxies does not matter so long as they are in the cluster and above the luminosity threshold. Observations of individual clusters with extensive spectroscopic data support the assumption that there is no substantial evolution in the relation between halo occupation number and cluster mass (Muzzin et al. 2007). This is also supported by several theoretical studies that find minimal evolution in the number of bright galaxies in massive halos (Kravtsov et al. 2004; Zentner et al. 2005).

We performed an independent validation of the MaxBCG relation with an analysis of the individual clusters in our sample with substantial membership data. While most of the low-redshift clusters have substantial membership data, these data generally do not extend to our estimate of  $R_{200}$  (Martini et al. 2007), nor is the X-ray coverage complete to this radius. Our spectroscopic coverage was often limited to the size of the *Chandra* field of view. However, two useful exceptions are Abell 89B and MS1008.1-1224 and in both cases estimates agree to within a factor of two. Our wide-field X-ray coverage of Abell 85 and Abell 754 (Sivakoff et al. 2008) were designed to sample a substantial fraction of the projected  $R_{200}$  and these values also agree well. Figure 2 illustrates the difference between the MaxBCG membership estimates and our spectroscopic estimates. The larger points have nearly complete spectroscopic coverage to  $R_{200}$ , while smaller points are substantially more incomplete. These points indicate that the error introduced by adopting the MaxBCG relation is approximately a factor of two. This error estimate is also consistent

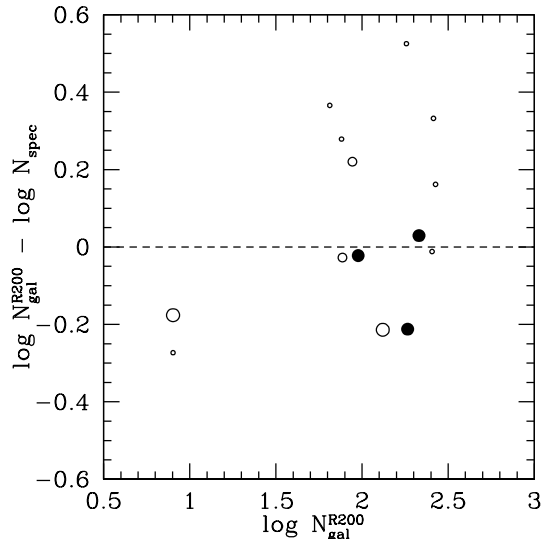


FIG. 2.— Difference between predicted and measured cluster richness compared to the cluster richness predicted by the MaxBCG sample. The quantity  $N_{gal}^{R200}$  is the number of red cluster galaxies more luminous than  $0.4L^*$  and estimated from the cluster velocity dispersion (Becker et al. 2007), while  $N_{spec}$  is a spectroscopic estimate of this quantity (see §4.3). Symbols are coded according to the spectroscopic completeness relative to  $R_{200}$ . Large circles have complete coverage to  $R_{200}$ , medium circles have more than 50% coverage, and the small circles have less than 50% coverage. Most clusters are at  $z < 0.5$  (open symbols), although substantial data exist for three at  $z > 0.5$  (filled symbols). See §4.3 for further details.

with an examination of figure 4 of Becker et al. (2007).

At higher redshifts three of our clusters have extensive membership information. We estimate that MS0015.9+1609 has  $\sim 200$  members based on several studies (Dressler & Gunn 1992; Ellingson et al. 1998) and that MS2053.7-0449 has  $\sim 100$  members (Tran et al. 2005). Note that these estimates are different from those presented in Eastman et al. (2007) due to updated completeness corrections and the change in the absolute magnitude threshold. For MS 1054-03 we estimate that there are  $\sim 300$  members from the extensive spectroscopic work of Tran et al. (2007). These three clusters are also shown in Figure 2 (filled circles). They are consistent with the low-redshift results and a factor of two uncertainty in the richness – velocity dispersion relation. While our estimates of the cluster galaxy population for these three clusters, as for the low-redshift clusters, are based on all galaxies rather than just red galaxies, the consistency supports the assumption that the integral of the bright end of the galaxy luminosity function in clusters above an evolving  $M_R$  threshold scales reasonably well with the cluster velocity dispersion independent of redshift, even if there is evolution in the colors of the cluster galaxies. The number of AGN, estimate of the inactive population, AGN fraction, and spectroscopic completeness for each cluster is listed in Table 7.

#### 4.4. Cluster AGN Fraction and Evolution

The AGN fraction for any single cluster is very small and it is uncertain due to small number statistics. In addition, the AGN fraction may vary from cluster to cluster due to correlations with other cluster properties such as velocity dispersion (Sivakoff et al. 2008). The AGN fraction may also depend on variations in the properties of the galaxy population within each cluster (e.g., mass, SFR, morphology). We therefore have binned the cluster sample in two ways to characterize variations with redshift. First, we simply split the sample

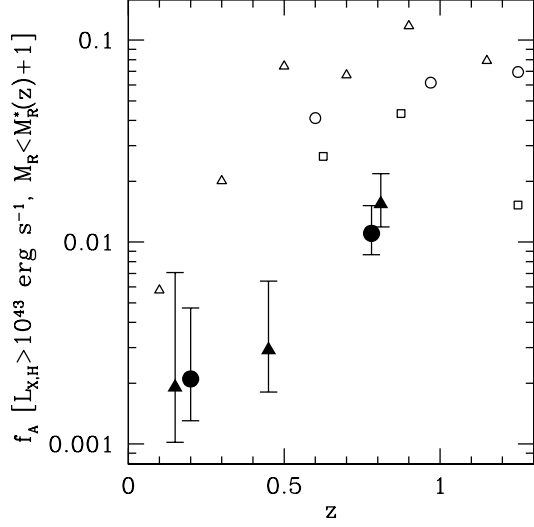


FIG. 3.— Evolution of the AGN population in clusters from  $z = 0$  to  $z = 1.3$  (filled symbols). The fraction of cluster members more luminous than  $M_R^* + 1$  with AGN that have  $L_{X,H} > 10^{43} \text{ erg s}^{-1}$  is shown in two redshift bins ( $z < 0.4$ ,  $z > 0.4$ ; filled circles) and three redshift bins ( $z < 0.3$ ,  $0.3 < z < 0.8$ ,  $0.8 < z < 1.3$ ; filled triangles). We also show our estimate of the field AGN fraction based on the galaxy LF estimates by Ilbert et al. (2005, open triangles), Dahlen et al. (2005, open circles), and Chen et al. (2003, open squares). See §4.4 for further details.

at  $z = 0.4$ . This choice is primarily motivated by the transition between where we rely on our own measurements and where we largely rely on other work. It also approximately divides the sample in two (17 clusters at  $z < 0.4$ , 15 at  $z > 0.4$ ). This yields completeness-corrected AGN fractions of  $f_A(z = 0.19) = 0.00134^{+0.0018}_{-0.00087}$  and  $f_A(z = 0.72) = 0.0100^{+0.0029}_{-0.0023}$ , or approximately a factor of eight increase in the AGN fraction (see Table 8) from a median redshift of 0.19 to a median redshift of 0.72. AGN fractions without the completeness correction are also listed in Table 8. The uncertainties on these quantities are double-sided,  $1 - \sigma$  confidence limits (Gehrels 1986). The increase in the AGN fraction is formally significant at the  $3.8\sigma$  level. We also split the sample into three bins with  $z < 0.3$ ,  $0.3 < z < 0.6$ , and  $z > 0.6$  to better resolve the continued increase at high redshift that is apparent in the raw data for individual clusters. This binning yields AGN fractions of  $f_A(z = 0.15) = 0.0019^{+0.0025}_{-0.0012}$ ,  $f_A(z = 0.45) = 0.0031^{+0.0024}_{-0.0015}$ , and  $f_A(z = 0.81) = 0.0147^{+0.0050}_{-0.0039}$ . The measured evolution between the lowest and highest bins is also a factor of eight and in good agreement with the other binning scheme. We note that the observed evolution is also well fit by a simple power law scaling as  $f_A \propto (1+z)^\alpha$  where  $\alpha = 5.3^{+1.8}_{-1.7}$ , although the power-law index is strongly correlated with the  $z = 0$  value of the AGN fraction.

The factor of eight evolution of the AGN fraction is smaller but consistent with the order of magnitude evolution observed by Eastman et al. (2007). They measured  $f_A(z = 0.2) = 0.0007^{+0.0021}_{-0.0007}$  and  $f_A(z = 0.6) = 0.020^{+0.012}_{-0.008}$  for  $L_{X,H} > 10^{43} \text{ erg s}^{-1}$ , although for a lower and fixed galaxy absolute magnitude of  $M_R = -20$ . At  $z = 0$  our galaxy absolute magnitude threshold is approximately a magnitude brighter than that used by Eastman et al. (2007) and the offset increases linearly with redshift. This difference in absolute magnitude threshold can readily account for the change in the low-redshift fraction because most of the AGN are associated with luminous cluster galaxies, that is increasing the galaxy luminosity threshold decreases the denominator and does not affect the numerator

TABLE 7  
AGN FRACTION ESTIMATES AND CLUSTER MEMBERSHIP

Cluster (1)	$z$ (2)	$\sigma$ (3)	$N_{AGN}$ (4)	$N_{gal}$ (5)	Flag (6)	$f_{A,raw} [\%]$ (7)	$f_{spec}$ (8)
Abell754	0.0546	953	1	82	1	$1.2^{+2.8}_{-1.0}$	1.00
Abell85	0.0554	993	0	53	1	$< 2.2$	1.00
Abell3128	0.0595	906	0	28	1	$< 4.1$	1.00
Abell3125	0.0616	475	0	15	1	$< 7.7$	1.00
Abell644	0.0701	952	0	40	1	$< 2.9$	1.00
Abell89B	0.0770	474	0	12	1	$< 9.6$	1.00
Abell2104	0.1544	1242	1	54	1	$1.9^{+4.3}_{-1.5}$	1.00
Abell1240	0.1590	698	0	28	2	$< 4.1$	1.00
Abell1689	0.1867	1400	0	184	1	$< 0.62$	1.00
Abell2163	0.2007	1381	0	262	1	$< 0.44$	1.00
Abell1942	0.2240	905	0	65	2	$< 1.8$	1.00
Abell2125	0.2465	1113	0	127	2	$< 0.90$	1.00
MS1455.0+2232	0.2578	1032	0	99	2	$< 1.2$	1.00
MS1008.1-1224	0.3068	1127	0	216	1	$< 0.53$	1.00
AC114	0.3148	1388	0	121	1	$< 0.95$	1.00
ZwCl1358.1+6245	0.328	1003	0	91	2	$< 1.3$	1.00
MS1512.4+3647	0.372	575	0	15	2	$< 7.7$	1.00
MS1621.5+2640	0.430	735	0	65	2	$< 1.8$	0.67
3C295	0.460	1642	2	412	2	$0.49^{+0.64}_{-0.31}$	0.67
MS0451.6-0305	0.538	1371	0	273	2	$< 0.42$	0.77
MS0015.9+1609	0.541	1234	1	214	2	$0.47^{+1.1}_{-0.39}$	0.77
RXJ0848.7+4456	0.574	895	1	102	2	$0.98^{+2.3}_{-0.81}$	0.67
MS2053.7-0449	0.583	865	0	95	2	$< 1.2$	1.00
RXJ0542.8-4100	0.634	1269	5	229	2	$2.18^{+1.5}_{-0.94}$	0.77
RXJ2302.8+0844	0.722	658	0	50	2	$< 2.3$	0.77
MS1137.5+6625	0.782	885	1	100	2	$1.00^{+2.3}_{-0.83}$	0.77
RX J1317.4+2911	0.805	1142	1	179	2	$0.56^{+1.3}_{-0.46}$	0.67
RXJ1716.4+6708	0.813	1445	3	308	2	$0.97^{+0.95}_{-0.53}$	0.92
MS 1054-03	0.823	1156	1	184	2	$0.54^{+1.2}_{-0.45}$	0.77
RDCS J0910+5422	1.110	675	1	53	2	$1.9^{+4.3}_{-1.6}$	0.67
Lynx E	1.261	740	1	66	2	$1.5^{+3.5}_{-1.3}$	0.67
Lynx W	1.270	650	1	49	2	$2.0^{+4.7}_{-1.7}$	0.67

NOTE. — AGN fraction estimates for individual clusters. Columns are: Col. (1): Cluster name; Col. (2): Redshift; Col. (3): Velocity dispersion (references for these values are in Table 1, Table 3, Sivakoff et al. (2008) for Abell 754, Abell 85, Abell 89B, Martini et al. (2006) for Abell 3128, Abell 3125, Abell 644, Abell 2104, Abell 2163, and MS1008.1-1224, or adopted from Czoske (2004) for Abell 1689 and Girardi & Mezzetti (2001) for AC 114); Col. (4): Number of AGN with  $L_{X,H} \geq 10^{43} \text{ erg s}^{-1}$  in galaxies more luminous than  $M_R^*(z) + 1$ ; Col. (5): Estimate of the number of cluster galaxies more luminous than  $M_R^*(z) + 1$  within either the *Chandra* FOV or  $R_{200}$ , whichever is smaller; Col. (6): Flag for the origin of the estimate where 1: from our spectroscopy and completeness correction; 2: from the MaxBCG as described in §4.3; Col. (7): Estimate of the cluster AGN fraction in percent; Col. (8): Estimate of the spectroscopic completeness for X-ray sources.

of the AGN fraction. In addition, we have since identified a second luminous AGN at low redshift (Sivakoff et al. 2008). At high redshift the change in galaxy luminosity threshold is also important, but in addition the cluster sample is more than three times larger than the Eastman et al. (2007) sample. The low-redshift cluster sample has increased by less than a factor of two.

One way to characterize the evolution of the cluster AGN fraction relative to the field is to calculate the integral of the field space density  $\Phi(L_{X,H} > 10^{43})$  as a function of redshift. Integration of the luminosity-dependent density evolution model in Ueda et al. (2003) yields a factor of five increase between  $z = 0.8$  and  $z = 0.2$ , which is somewhat less but consistent with the observed evolution of cluster AGN. However, this is not a fair comparison because the evolution of field AGN with  $\Phi(L_{X,H} > 10^{43})$  is not normalized by the evolution of all field galaxies brighter than  $M_R^* + 1$  and the cluster AGN fraction is.

While there is not a direct measurement of the field AGN fraction similar to our calculation for clusters (although see Lehmer et al. 2007), we can estimate this quantity by dividing the integral of the field hard X-ray LF from Ueda et al. (2003)

TABLE 8  
AGN FRACTION FOR SUBSAMPLES OF THE CLUSTERS

Sample (1)	$z$ range (2)	$N_{CL}$ (3)	median $z$ (4)	median $\sigma$ (5)	$N_{A,raw}$ (6)	$N_{gal}$ (7)	$f_{A,raw}$ [%] (8)	$f_{spec}$ (9)	$f_{A,corr}$ [%] (10)
Two Bins									
	$z < 0.4$	17	0.19	993	2	1492	$0.134^{+0.18}_{-0.087}$	1.00	$0.134^{+0.18}_{-0.087}$
	$z > 0.4$	15	0.72	895	18	2379	$0.76^{+0.22}_{-0.18}$	0.76	$1.00^{+0.29}_{-0.23}$
Three Bins									
	$z < 0.3$	13	0.15	953	2	1049	$0.19^{+0.25}_{-0.12}$	1.00	$0.19^{+0.25}_{-0.12}$
	$0.3 < z < 0.6$	10	0.45	1065	4	1604	$0.25^{+0.20}_{-0.12}$	0.81	$0.31^{+0.24}_{-0.15}$
	$z > 0.6$	9	0.81	885	14	1218	$1.15^{+0.42}_{-0.39}$	0.78	$1.47^{+0.45}_{-0.39}$

NOTE. — Cluster AGN fractions with the data split into two bins and three bins. The two bins are split at  $z = 0.4$ , while the three bins split the data at  $z = 0.3$  and  $z = 0.6$ . For each bin we list: Col. (2): redshift range; Col. (3): number of clusters; Col. (4): median redshift; Col. (5): median velocity dispersion of clusters; Col. (6): sum of the luminous AGN in the bin; Col. (7): raw AGN fraction with double-sided,  $1 - \sigma$  confidence limits; Col. (8): estimate of the mean spectroscopic completeness weighted by the number of galaxies per cluster; Col. (9): AGN fraction corrected for spectroscopic completeness.

by the integral of the galaxy LF. We have identified three surveys that report LF measurements for the  $R$ -band and approximately span the same redshift range of this work. The first of these is the VIMOS-VLT Deep Survey (Ilbert et al. 2005), which is based on  $UBVRI$  photometry,  $\sim 11,000$  spectra to  $I_{AB} = 24$  mag and extends from  $z = 0.05$  to  $z = 2$  (although their lowest-redshift point is taken from SDSS; Blanton et al. 2003). We also show results from two measurements based on photometric redshift data: the Las Campanas Infrared Survey (LCIRS; Chen et al. 2003), which is mostly based on  $UBVRIH$  measurements and presents the LF for  $z = 0.5 - 1.5$ , and the Great Observatories Origins Deep Survey (GOODS; Dahlen et al. 2005), which is based on  $U$  through  $K$  observations and presents the galaxy LF to  $z = 2$ . While these photometric redshift surveys may have more systematic uncertainties than the LF based on spectroscopic measurements, they have the virtue that they have measured the luminosity function in the rest-frame  $R$ -band, rather than relied on assumptions about galaxy spectral energy distributions (SEDs) to calculate  $k$ -corrections. We have calculated the field AGN fraction for each of these surveys and show the results in Figure 3 (*open symbols*). At low redshift the AGN fraction calculated with the Ilbert et al. (2005) LF is approximately a factor of five above the cluster fraction, which is consistent with the difference between the field and clusters seen by Dressler et al. (1999) for spectroscopically-identified AGN. At higher redshifts ( $z > 0.5$ ), the field estimates range between a factor of three and a factor of ten above the cluster fraction. These estimates of the field AGN fraction vary so substantially due to the dispersion in estimates of the galaxy luminosity function. In addition, this calculation presupposes that all of the X-ray AGN are in galaxies more luminous than  $M_R^*(z) + 1$ . While there is good evidence that most of these luminous X-ray AGN are in relatively luminous galaxies (e.g. Silverman et al. 2009a), there is nevertheless a bias against spectroscopic identification of lower-luminosity X-ray AGN host galaxies. Finally, we note that the relative evolution of galaxies in clusters and the field further complicates this comparison. In future work we hope to compile sufficient data to calculate the AGN fraction in the field and clusters as a function of galaxy mass. At present the data are insufficient to conclude if the cluster AGN fraction or field AGN fraction evolves more rapidly.

#### 4.5. Contamination by AGN Associated with Large-Scale Structure

One concern raised about the physical origin of the Butcher-Oemler effect is the contribution of projection effects. Diaferio et al. (2001) studied this issue in detail with N-body simulations and semianalytic models to distinguish true cluster members from field interlopers that were at the cluster redshift and within the projected  $R_{200}$ , yet physically outside the cluster  $R_{200}$ . Diaferio et al. (2001) concluded that up to 50% of the apparent Butcher-Oemler galaxies at the redshifts of high-redshift clusters may be interlopers. A similar effect may be relevant for the AGN population and such a large contamination would decrease the observed evolution, but not erase it.

While there is no comparable study that directly investigates the projection of AGN onto high-redshift clusters, there is good evidence that AGN are associated with the large-scale environment of clusters. Gilmour et al. (2007) identified 11 X-ray AGN (to a lower luminosity limit of  $\sim 10^{41}$  erg s $^{-1}$ ) in the A901/2 supercluster at  $z \sim 0.17$  and only one was in the densest region of the supercluster. The remainder were mainly in regions of intermediate density. In the vicinity of 3C295 ( $z = 0.46$ ) D’Elia et al. (2008) find evidence for AGN associated with a filamentary structure. At yet higher redshifts this trend is also apparent. Kocevski et al. (2009) find X-ray AGN associated with the CL1604 supercluster at  $z \sim 0.9$ , which contains 8 confirmed groups and clusters. These AGN mostly avoid the densest regions of the clusters and are located on the outskirts of the most massive clusters, that is they are associated with poorer clusters and groups.

We examined our data to determine if there were a population of AGN outside the projected  $R_{200}$  for these clusters similar to those seen in the two superclusters. This is only possible with the subset of the sample with substantial coverage beyond  $R_{200}$ . Eight of the clusters have *Chandra* coverage that extends to  $2R_{200}$ . There are six AGN between  $R_{200}$  and  $2R_{200}$  that meet our velocity cuts for cluster membership compared to eight AGN within  $R_{200}$  for these same clusters. The larger number within the clusters suggests the opposite trend from the two supercluster studies described above, although these results are not truly in conflict because the supercluster studies encompassed a much larger area outside of dense clusters than this study. The different large-scale environments associated with these clusters and the superclusters suggest a more quantitative comparison would not be meaningful. These large-scale structure data also provide a crude means to estimate the likelihood of chance juxtapositions of AGN associated with large-scale structure onto the clusters. If interloper AGN have the same surface density within  $R_{200}$

as between  $R_{200}$  and  $2R_{200}$ , then the six we identified in an area of  $3\pi R_{200}^2$  suggest we should expect at most 2 interlopers compared to the 8 AGN we see within  $R_{200}$ . This line of argument suggests that the interloper fraction is 25%, which is small compared to the observed evolution signature.

#### 4.6. Uncertainties

One major potential source of systematic error is the use of the MaxBCG richness estimator to estimate the fraction of cluster galaxies more luminous than  $M_R^* + 1$ . In §4.3 we estimated that there is a factor of two uncertainty in the use of this relation. This uncertainty is mainly important for the high-redshift subsamples as the low-redshift subsamples have more complete spectroscopic membership data. If we randomly introduce a factor of two uncertainty in each cluster, the effect is negligible when averaged over the 15 clusters with  $z > 0.4$  compared to the factor of eight evolution in the AGN fraction.

As mentioned previously, another valid concern with the MaxBCG estimator is that it is calibrated to the number of red galaxies in the cluster and this population may not all be in place at  $z = 0.4$  and higher. For our application it does not matter if the galaxies are red or not, just that they are in the cluster. Furthermore, if we have overestimated the number of galaxies brighter than  $M_R^* + 1$  then we have underestimated the evolution of the AGN fraction and our result is yet more statistically significant. The assumption that all of the galaxies are red does impact the  $k$ -correction we use to estimate the spectroscopic limit corresponding to  $M_R^*(z) + 1$  and thus the size of our completeness correction. If the galaxies are redder, then the  $k$ -correction would be smaller, the apparent magnitude limit would be brighter, and the completeness correction would be smaller. The implication would be that we have preferentially overestimated the AGN fraction at high redshifts because completeness corrections are only applied to the high-redshift clusters. While the average completeness correction approaches 25% (see Table 8), in practice the spectroscopic completeness is not a strong function of apparent magnitude (e.g. see §4.2, Silverman et al. 2005; Eckart et al. 2006) and we consequently expect much less than a 25% reduction in the evolution. The evolution of the host galaxy population is also important because if there were less fading of  $M^*(z)$  than we assume, then the completeness limit would be too bright and we would have underestimated the AGN fraction at high redshift.

The value of the cluster velocity dispersion introduces additional uncertainty to this calculation in two ways. First, many of the direct measurements of the cluster velocity dispersion, particularly for high-redshift clusters, are based on small samples of galaxies and thus the velocity dispersion itself may be uncertain, particularly if the galaxy velocity distribution is not Gaussian. Second, as noted above the cluster velocity dispersion has not been directly measured for several clusters and we instead used the X-ray temperature and the results of Xue & Wu (2000) to estimate the velocity dispersion and this has a 30% scatter. We checked both of these concerns with a measurement of the scatter between  $\sigma$  and  $T_X$  for the ten high-redshift clusters with measurements of both quantities and the mean deviation is  $\sim 220 \text{ km s}^{-1}$  if we exclude 3C295, which has a substantially higher velocity dispersion (1642 km/s Girardi & Mezzetti 2001) than expected from its X-ray temperature (5.3 K from Vikhlinin et al. 2002). This mean deviation corresponds to approximately a factor of two uncertainty in the richness, which is comparable to the uncertainty we derived for the richness estimator. From this analy-

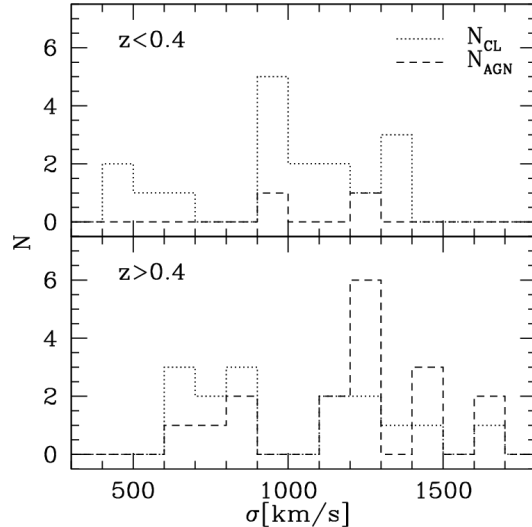


FIG. 4.— Histograms of the number of clusters with a given velocity dispersion (dotted line) and the number of AGN in clusters of a given velocity dispersion (dashed line) for the low-redshift ( $z < 0.4$ ; top panel) and the high-redshift ( $z > 0.4$ ; bottom panel) subsamples. The cluster samples are reasonably well matched within these two redshift bins.

sis we similarly conclude that this source of uncertainty does not substantially affect our results.

A related evolutionary effect is that the velocity distributions of the high-redshift clusters may be systematically more non-Gaussian than low-redshift clusters because the high-redshift clusters are less likely to be relaxed. If the cluster velocity dispersion were overestimated, then the richness and  $R_{200}$  would be overestimated as well. This in turn would lead to an underestimate of the AGN fraction in high-redshift clusters. Jeltema et al. (2005) measured power ratios from *Chandra* observations of the IGM for a large sample of clusters out to  $z \sim 1$  and found good evidence that high-redshift clusters are less relaxed than low-redshift clusters, so this potential source of systematic error would lead us to underestimate the AGN fraction. Nine of our clusters were analyzed in the Jeltema et al. (2005) study, including eight in our high-redshift sample. We compared the AGN fractions and the power ratios for these clusters, but did not find a significant trend. Unfortunately we do not have sufficient redshift data for most high-redshift clusters to look for non-Gaussianity in the galaxy velocity distribution, although note there is no evidence for a trend between dynamically-disturbed clusters and AGN fraction at low redshift (Martini et al. 2007).

Finally, we consider the evolution of the cluster population to determine if the higher-redshift clusters represent the progenitor population of the lower redshift clusters. As noted previously, observations at low redshift indicate that the AGN fraction depends on environment and specifically that the AGN fraction is higher in lower velocity dispersion environments (Sivakoff et al. 2008; Arnold et al. 2009). Therefore if our high-redshift clusters are the progenitors of lower velocity dispersion clusters or massive groups, then the observed evolution may not be as significant. As many of the high-redshift clusters are X-ray selected, they are generally high-mass clusters and are reasonably well matched to the lower-redshift sample (see Figure 4 and Table 7). Following Finn et al. (2005) and Poggianti et al. (2006), we have estimated the velocity dispersions of the progenitors of the high-redshift cluster sample and find they are in good agreement. For exam-

TABLE 9  
HIGH-REDSHIFT AGN ASSOCIATED WITH LARGE-SCALE STRUCTURE AROUND CLUSTERS

AGN (1)	Cluster (2)	$z$ (3)	$R$ [mag] (4)	$\log L_{X,H}$ [erg s $^{-1}$ ] (5)	$\delta v/\sigma$ [km/s] (6)	$\Delta R$ [arcmin] (7)	$R/R_{200}$ (8)	Class (9)
CXOSEXSI J084846.0+445945	RX J0848.7+4456	0.567	21.45	43.1	1.99	3.51	1.16	ELG
CXOMP J230300.9+084659	RXJ2302.8+0844	0.738	21.71	44.23	2.81	4.46	1.2	BLAGN
CXOSEXSI J171807.6+670647	RXJ1716.4+6708	0.797	21.75	44	1.83	7.8	1.59	BLAGN
CXOU J105710.6-033500	MS 1054-03	0.832	21.93	43.14	1.27	4.57	1.18	ALG
CXOSEXSI J091040.8+542006	RDCS J0910+5422	1.097	22.38	43.1	2.74	2	1.13	ELG
CXOSEXSI J084903.9+445023	LynxE	1.276	23.92	43.2	2.95	1.76	1.03	ELG

NOTE. — AGN associated with large-scale structure around the subset of high-redshift clusters with complete X-ray coverage to twice the projected virial radius. This is the subset of AGN that satisfy the redshift selection criterion, but have a projected distance of  $1 < R/R_{200} \leq 2$ . Columns are identical to Table 2. The data for CXOU J105710.6-033500 are from van Dokkum et al. (2000) for the redshift, magnitude, and classification and the X-ray data are from Johnson et al. (2003). This sample is described in further detail in §4.5.

ple, the progenitor of a 1000 km s $^{-1}$  cluster at  $z = 0$  has 800 km s $^{-1}$  at  $z = 0.6$  (Poggianti et al. 2006), or only about 100 km s $^{-1}$  less than the difference between our low-redshift and high-redshift subsamples. The sense of this trend is that the high-redshift sample is actually somewhat more massive than the typical progenitor of the low-redshift sample and therefore the minor mismatch in cluster masses is more likely to have dampened rather than enhanced the measured evolution of the AGN fraction.

## 5. PROPERTIES OF THE CLUSTER AGN

### 5.1. Distribution

The projected radial and velocity distributions of the AGN provide valuable additional information about the origin of the AGN. For example, if the AGN are preferentially located in the cluster outskirts, or preferentially have a higher velocity dispersion than the cluster mean, this may indicate that their host galaxies have relatively recently entered the cluster potential. This is known to be the case for emission-line galaxies (Biviano et al. 1997; Dressler et al. 1999). At low-redshifts and for lower-luminosity X-ray AGN, Martini et al. (2007) found that  $L_X > 10^{42}$  erg s $^{-1}$  [0.5–8 keV] AGN were more centrally concentrated than typical cluster galaxies, while AGN an order of magnitude less luminous had the same distribution as the inactive galaxy population. For both luminosity thresholds the velocity distribution of the AGN were consistent with the galaxy population.

It is more challenging to compare these higher-luminosity X-ray AGN to the host galaxy population because we lack membership data for nearly all of the high-redshift clusters. Nevertheless, we can compare the distribution of sources to the typical distribution of cluster galaxies and to the excess surface density distribution found by surveys of X-ray point sources toward distant clusters. In Figure 5 we present a histogram of the number of X-ray AGN from the cluster center as a function of distance in both physical units (Mpc) and normalized to  $R_{200}$ . While the sample is small, two results are apparent from the figure. First, there are approximately equal numbers of AGN outside  $0.5R_{200}$  as inside it, whereas if the AGN traced the cluster galaxy distribution we would expect them to be more centrally concentrated. Second, the radial distribution is more strongly peaked when plotted in physical units than normalized to  $R_{200}$ .

While we do not have detailed information on the radial distribution of the cluster galaxy populations in these clusters, we do have extensive data on nearby clusters from Christlein & Zabludoff (2003). For these clusters we have investigated the cluster galaxy distribution with the same selection criteria ( $R < R_{200}$ ,  $M_R < M_R^* + 1$ ,  $\Delta v < 3\sigma$ ) and find that 70% of the galaxies fall within  $0.5R_{200}$ , whereas 10 of

18 luminous AGN at  $z > 0.4$  are within  $0.5R_{200}$ . The binomial probability is only 14% that we would find 10 or fewer AGN within  $0.5R_{200}$  if we expected 70%. There is thus a mild tendency for luminous AGN to be distributed toward the outskirts of the clusters, although this does make the substantial assumption that the radial distribution of galaxies within clusters is similar at  $z \sim 0.8$  and the present. This broad distribution in radius is in contrast to our earlier results on lower-luminosity AGN in lower-redshift clusters. At low redshift we found that 50% of the luminous AGN were within  $0.1R_{200}$  (Martini et al. 2007). Better statistics could determine if the AGN are preferentially located in the outskirts of clusters compared to all cluster galaxies. That would be consistent with the hypothesis that AGN are triggered by mergers during infall. From simulations Ghigna et al. (1998) find that mergers between galaxies do not occur within the virial radius. We note that Berrier et al. (2009) simulated the formation of 53 galaxy clusters and find most cluster galaxies do not experience ‘preprocessing’ in group environments and therefore processes specific to clusters must largely be responsible for the differences between cluster and field galaxies.

The second result has interesting implications for studies that use the surface density distribution of excess sources to characterize the distribution of AGN in clusters (Ruderman & Ebeling 2005; Gilmour et al. 2009; Galametz et al. 2009). These studies generally plot the excess surface density as a function of physical distance from the cluster center and find a central peak in surface density. Our results indicate that the true distribution may be flatter than implied by use of the physical (proper) distance from the cluster core. This is because those surveys, like the present study, include clusters with a wide range of masses and consequently a wide range of  $R_{200}$ . Simply adding the distributions for all clusters without renormalizing each observation for the size of the cluster will produce an artificial central peak due to the mass range of the cluster sample.

If the cluster AGN are associated with a population that recently entered the cluster potential, the host galaxies may also be preferentially on more radial orbits and have a larger velocity dispersion than that of all cluster galaxies. As noted previously, this is true of the emission-line galaxy population in clusters. In Figure 6 we plot the cumulative velocity distribution for all 18 AGN with  $z > 0.4$  normalized by the cluster velocity dispersion. The distribution is in excellent agreement with a Gaussian distribution and we therefore find no evidence that the cluster AGN have a larger velocity distribution that would be consistent with more radial orbits. This was also found for the 14 relatively luminous ( $L_{X,B} \geq 10^{42}$  erg s $^{-1}$ ) AGN studied by Martini et al. (2007). A better test would be to compare the AGN host population to the absorption-line

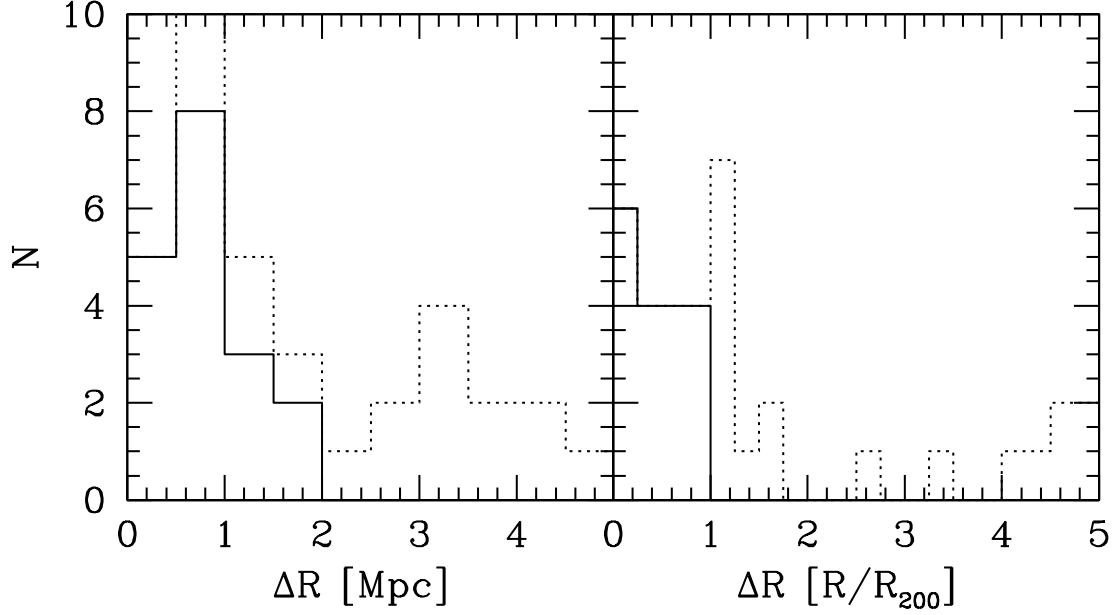


FIG. 5.— Histograms of the AGN clustercentric distances in terms of Mpc (*left*) and normalized to  $R_{200}$  (*right*) for cluster AGN with  $z > 0.4$ . The distribution of the confirmed cluster members (*solid line*) is much more centrally peaked when expressed in terms of Mpc than in terms of  $r/R_{200}$ . Other AGN associated with large-scale structure (with  $R > R_{200}$ ) are also shown (*dotted line*).

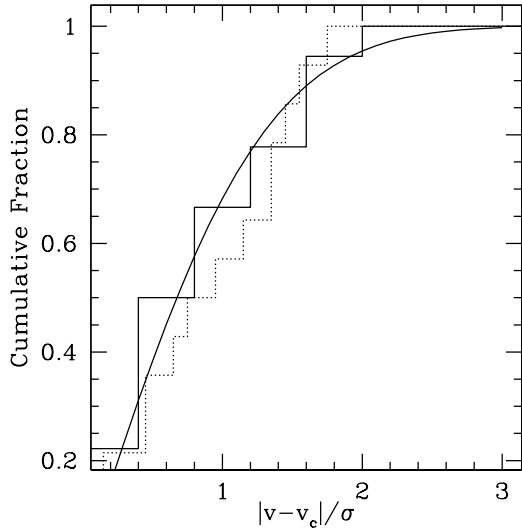


FIG. 6.— Histogram of the cumulative velocity distribution of cluster AGN normalized to the cluster velocity dispersion for the 18 cluster AGN with  $z > 0.4$  (*solid histogram*). The AGN velocity distribution is consistent with a Gaussian distribution (*solid curve*) and the  $L_{X,B} \geq 10^{42}$  erg s $^{-1}$  AGN from Martini et al. (2007) (*dotted histogram*).

galaxies in the clusters since the velocity dispersion estimates for many of these clusters may be biased toward the emission-line galaxy population because it is easier to measure redshifts for them. While this is not the case for those whose velocity dispersions are estimated from X-ray data, it may also be true of the calibration sample for the relations between X-ray properties and galaxy velocity dispersion.

### 5.2. Luminosity Function

We have begun to acquire sufficient numbers of cluster AGN that it is possible to compare the X-ray luminosity function (XLF) between clusters and the field, as well as the cluster XLF at different redshifts. A comparison between the

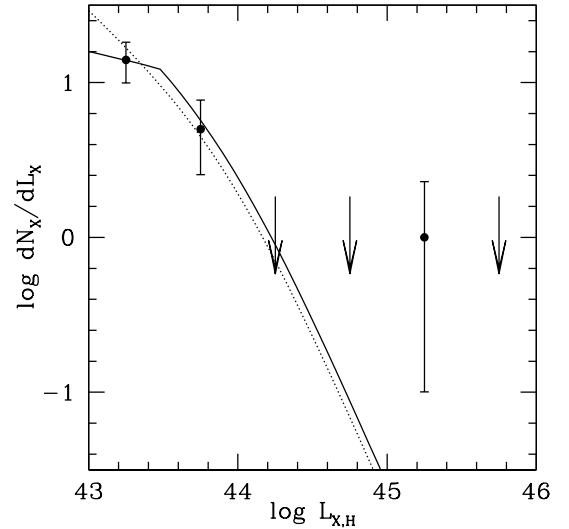


FIG. 7.— Hard X-ray luminosity function of cluster AGN at  $z > 0.4$  compared to the field XLF from Ueda et al. (2003) at the median cluster redshift ( $z = 0.8$ , *solid curve*) and at low redshift ( $z = 0.1$ , *dotted curve*). The field XLFs have been renormalized to be consistent with the cluster measurements in the first two luminosity bins. The arrows are upper limits calculated with Poisson statistics.

cluster and field XLF is interesting because differences between the two would be a signature of environment-dependent downsizing. There is evidence that this is true of star formation in different environments. For example Kauffmann et al. (2004) find that substantial star formation is only present in higher mass galaxies in lower density environments in the local universe. If the cluster black holes primarily grew at higher redshifts than field black holes, similar to the earlier formation epoch expected for the stellar populations in luminous cluster galaxies, then the cluster luminosity function at high-redshift may have a similar shape to the present-day XLF in the field. One test of this hypothesis is to compare the characteristic luminosity  $L_X^*$  between clusters and the field. If the

cluster AGN primarily grew at an earlier epoch,  $L_X^*$  would be smaller in clusters relative to the field at a given redshift.

It is reasonable to compare the shape of the XLF between clusters and the field because the XLF is a measurement of the X-ray sources alone within well-defined volumes, although the caveats associated with large scale structure discussed in §4.5 do apply. This is different from the case in Section 4.4, where we noted that the comparison of the evolution of the AGN fraction and the integrated space density was not comparing identical quantities because the AGN fraction includes information about the galaxy population. The one assumption that we do make is that all of the X-ray sources are hosted by galaxies above our threshold, but this is reasonable given Figure 1. In addition, the normalization remains arbitrary because it is challenging to define a total volume for the cluster AGN sample, although this is not necessary because the shape of the XLF already provides useful information. In Figure 7 we plot the cluster XLF for our  $z > 0.4$  sample compared to the field XLF at the median cluster redshift of  $z = 0.8$  from Ueda et al. (2003). The cluster XLF is in reasonable agreement with the field XLF at the same redshift, although the statistics are quite limited. As motivation for future work, we also plot the field XLF at lower redshift ( $z = 0.1$ , *dotted line*). For the lower-redshift XLF  $L_X^*$  is smaller and consequently all  $L_{X,H} \geq 10^{43}$  AGN are above the characteristic luminosity, while these data straddle  $L_X^*$  in the field XLF at  $z = 0.8$ . Improved statistics for cluster X-ray AGN at  $z > 0.4$  could determine if there is also a break in the cluster XLF, or if it is more similar to the field XLF at lower redshift.

The evolution of the cluster XLF with redshift is also relevant for the origin of X-ray AGN in lower-redshift clusters. If cluster AGN at the present day are simply the descendants of AGN at higher redshift that have been fading for several Gyr, then the difference between the low-redshift and high-redshift cluster XLF should be consistent with pure luminosity evolution. In contrast, if there is substantial retriggering of low-luminosity AGN in low-redshift clusters, or if other mechanisms are capable of fueling AGN in clusters, then the cluster XLF evolution may not be consistent with just luminosity evolution. A signature of other fueling or triggering mechanisms would be a substantially larger population of lower-luminosity AGN in present day clusters compared to expectations from the high-redshift population. While pure luminosity evolution would be surprising because this is not observed in field AGN, the most luminous cluster galaxies are consistent with passive evolution. Better measurements of the cluster XLF over a broader range in luminosity could investigate this hypothesis.

### 5.3. Host Galaxy Properties

Both the colors and morphologies of low-luminosity ( $\sim 10^{41}$  erg s $^{-1}$ ) AGN in low-redshift clusters suggest they are primarily hosted by galaxies dominated by light from their old stellar populations (Martini et al. 2006). This becomes progressively less true for higher-luminosity AGN and ground-based observations of the most luminous sources ( $\geq 10^{43}$  erg/s) in Abell 2104 (Martini et al. 2002) and Abell 754 (Arnold et al. 2009) indicate that they have late-type morphologies, although their hosts are luminous. In addition, these more luminous AGN are more likely to exhibit visible-wavelength AGN spectral signatures than their lower-luminosity counterparts.

While the spectroscopic classification of the high-redshift sample is fairly subjective because of variations in wavelength

coverage and signal to noise ratio, the spectroscopic classifications reported by Silverman et al. (2005) and Eckart et al. (2006) support the low-redshift results. They classified six of 17 X-ray AGN as BLAGN, nine as other emission-line galaxies, and the remaining two as absorption-line galaxies. The vast majority of the higher luminosity AGN have substantial line emission, even with the bias against redshift measurements for sources without strong emission lines. We note for comparison that two of the six AGN in the large-scale structure sample are classified as BLAGN and the other four are evenly split between emission-line and absorption-line galaxies. These other sources are thus similar to the cluster AGN.

Several of the high-redshift clusters also have HST observations suitable to study the morphologies of the cluster galaxies. The largest survey of X-ray source morphology in high-redshift clusters is that by Martel et al. (2007), who investigate the fields of five high-redshift clusters: RX J0152-1357, RX J0849+4452, RDCS J0910+5422, MS 1054-0321, and RDCS J1252-2927, and the middle three clusters overlap this sample. For the entire field sample they classify half of the X-ray counterparts as early-type, 35% as late-type, and 15% as irregular galaxies. For the six cluster members in their sample, they find half are in early-type hosts, two in late-type hosts, and one in an irregular galaxy. In addition, three of these cluster AGN hosts are in interacting systems. The specific overlap with this sample are CXOU J091043.3+542152 and CXOMP J105650.6-033508 and both have early-type morphologies (their other member in RDCS J0910+5422 falls slightly below our luminosity threshold).

### 5.4. Implications for the Sunyaev-Zel'dovich Effect

Many cluster surveys are currently planned or in progress that use the Sunyaev-Zel'dovich effect to identify large numbers of clusters (e.g. Kosowsky 2003; Ruhl et al. 2004). This effect is caused by inverse Compton scattering of Cosmic Microwave Background (CMB) photons off hot electrons in the ICM that changes the spectrum of the CMB in the direction of a cluster (e.g. Carlstrom et al. 2002). The main virtue of this effect is that it is redshift independent, and consequently can be used to detect (the hot electrons associated with) clusters out to high redshifts. However, mechanical heating by AGN in the cluster may contribute to the thermal energy of the ICM (e.g. Bîrzan et al. 2004) and thus make it more difficult to identify some clusters. Any increase in the AGN population with redshift will also introduce a systematic effect with redshift.

The potential impact of AGN on SZE cluster surveys was recently examined in detail by Lin & Mohr (2007). They measured the radio luminosity function in nearby clusters at 1.4 GHz and used measurements of AGN at higher frequencies (Cooray et al. 1998; Coble et al. 2007) to estimate that on order 10% of clusters will have an AGN flux comparable to the SZE flux. As a worst-case scenario they adopted an evolution model where the fraction of radio AGN increases as  $(1+z)^{2.5}$ . This model was largely motivated by observations of the radio galaxy luminosity function, which suggested evidence for an increase (Best et al. 2002; Branchesi et al. 2006). If this population evolves at a comparable rate more consistent with the  $(1+z)^{5.3}$  rate we observe for luminous X-ray AGN, then the fraction of substantially contaminated clusters will be higher than predicted by Lin & Mohr (2007).

## 6. DISCUSSION



The extent of the correlation between the evolution of star formation and AGN in clusters could provide valuable new insights into how closely related these two processes are. The original work by Butcher & Oemler (1978, 1984) on the evolution of the fraction of blue galaxies in clusters provides a useful first point of comparison to the AGN fraction evolution, in part because we adopted many elements of their methodology. Specifically, Butcher & Oemler (1984) characterized cluster galaxy evolution with: 1) a fixed criterion to define the sample of interest (a galaxy was classified as blue if the rest-frame  $B - V$  color was at least 0.2 mag bluer than the relation exhibited by the red galaxies); 2) measurement of this population relative only to cluster galaxies above some luminosity threshold ( $M_V = -20$ ); 3) use of an aperture scaled to the physical properties of individual clusters (a circle that contained the inner 30% of the cluster galaxy population). With these definitions, Butcher & Oemler (1984) found that the blue galaxy fraction increased from  $f_B \sim 0.03$  at  $z \leq 0.1$  to  $f_B \sim 0.25$  at  $z = 0.5$  for relatively compact, concentrated clusters, or approximately an order of magnitude.

One of the most recent and comprehensive studies of the evolution of star formation in clusters is the work of Poggianti et al. (2006). These authors used the  $[\text{O II}]\lambda 3727$  line as a tracer of star formation, rather than color, and measured the fraction of galaxies with  $[\text{O II}]$  emission (equivalent width  $> 3\text{\AA}$ ) as a function of both cluster redshift and cluster velocity dispersion. Their sample includes 25 clusters with  $z = 0.4 - 0.8$  and another 10 groups in the same redshift range, while they have a large local comparison sample at  $z = 0.04 - 0.08$  from the Sloan Digital Sky Survey. They measure the  $[\text{O II}]$  fraction  $f_{[\text{O II}]}$  relative to an evolving absolute magnitude limit  $M_{V,\text{lim}}$  that varies from  $-20.5$  at  $z = 0.8$  to  $-20.1$  at  $z = 0.4$ , while the local limit was  $M_V < -19.8$ . Their main results are that there is substantial evolution in  $f_{[\text{O II}]}$  and that there is substantial variation in  $f_{[\text{O II}]}$  with velocity dispersion at a given redshift. Given the velocity dispersion dependence, a direct comparison of the evolution of  $f_{[\text{O II}]}$  with  $f_A$  is not meaningful for different cluster samples. Instead, we have used their upper envelope for  $f_{[\text{O II}]}(\sigma)$  at high redshift and their envelope prescription at low redshift to estimate  $f_{[\text{O II}]}$  for each of our clusters and then computed the average  $f_{[\text{O II}]}$  for each of the subsamples shown in Table 8. These relations predict an increase in  $f_{[\text{O II}]}$  of less than a factor of two from the low-redshift to the high-redshift subsamples, or substantially less than the factor of eight we observe for the AGN fraction.

These results are interesting, although numerous caveats forestall too much interpretation of the relative rates of evolution. One major concern is that there is likely downsizing in clusters similar to what is observed in the field (e.g. Cowie et al. 1996; Hasinger et al. 2005; Silverman et al. 2008; Yenko et al. 2009), that is the relative number of galaxies with star formation or AGN activity above a certain threshold varies with redshift. The direct implication of this for the AGN fraction is that the evolution of the AGN fraction over a given redshift range is expected to depend on luminosity, just as the rate of evolution of the AGN space density is observed to vary in the field as a function of minimum luminosity (e.g. Ueda et al. 2003). This is similarly a complication for interpretation of the evolution of star formation, and consequently limits direct comparison of the mere rates of evolution of star formation and AGN above some threshold. For example, while Poggianti et al. (2006) have simi-

larly used an evolving galaxy luminosity threshold to characterize the evolution of the star-forming galaxy fraction, their galaxy luminosity threshold is over a magnitude fainter and therefore they have measured the evolution in a population that includes many more fainter cluster members. However, these concerns are not an obvious limitation to comparisons that use the same luminosity threshold to separately compare either AGN or star formation across different environments, particularly when the evolution of the star formation rate and AGN luminosity are tied to the same galaxy population. For example, if the relative rates of evolution of AGN and star formation in  $< M_R^* + 1$  galaxies were different in the field and clusters, this would suggest a limit to the extent of the apparent coevolution of black holes and galaxies in at least one of these environments.

Another concern about a direct comparison to these measurements of the evolution of the star forming galaxy population is that  $[\text{O II}]$  emission is more susceptible to reddening and metallicity effects relative to other star formation indicators, such as  $\text{H}\alpha$  (Kewley et al. 2004). Many ISO studies (summarized by Metcalfe et al. 2005) found evidence for an increase in star formation in clusters at higher redshifts, and that the increase appeared to be greater than that predicted by UV continuum or visible-wavelength spectroscopic diagnostics. *Spitzer* observations of clusters have also found substantial, often obscured, star formation in high-redshift clusters (Geach et al. 2006; Marcillac et al. 2007; Bai et al. 2007). Geach et al. (2006) used new *Spitzer* data for two clusters and data for five others from the literature to estimate the star formation rate normalized by the cluster mass. They find evidence for an increase in higher redshift clusters, but also substantial variation between clusters at the same redshift. Saintonge et al. (2008) used a larger sample of eight clusters with  $24\mu\text{m}$  *Spitzer* data to study the evolution of the fraction of obscured star-forming galaxies from  $z = 0.02 - 0.83$ . They find that the fraction of cluster galaxies with star formation rates above  $5 M_\odot \text{yr}^{-1}$  increases from 3% at  $z = 0.02$  to 13% at  $z = 0.83$  and that this is stronger evolution than exhibited by color-selection, such as the criteria of Butcher & Oemler (1978, 1984). The star-forming galaxies they identify in these clusters are also mostly disjoint from the Butcher-Oemler galaxies and consequently when they sum the blue and mid-infrared galaxies the fraction of star-forming galaxies increases to  $\sim 23\%$  at high redshift.

Several of these *Spitzer* studies overlap clusters that are also in our sample and it is interesting to see if there is a direct correspondence between the AGN and mid-infrared sources detected by *Spitzer*. The massive cluster MS1054-03 was studied by Bai et al. (2007) and their  $24\mu\text{m}$  sources include the two X-ray AGN identified by Johnson et al. (2003). Saintonge et al. (2008) have three clusters in common with our sample: MS0451.6-0305, MS2053.7-0449, and MS 1054-03, although they do not provide information on individual sources. While not in our sample, the study of RX J0152.7-1357 ( $z = 0.831$ ) by Marcillac et al. (2007) found that the two most luminous  $24\mu\text{m}$  sources (of 22 confirmed members) were also X-ray AGN. Similarly, Geach et al. (2009) found that one (of 12) of the luminous infrared galaxies ( $L_{\text{IR}} > 10^{11} L_\odot$ ) in CL0024+16 ( $z = 0.4$ ) was obviously an AGN based on their infrared data alone. At lower redshifts, Gallagher et al. (2008) have also used *Spitzer* data to identify AGN and star forming galaxies in Hickson Compact Groups.

Saintonge et al. (2008) explore whether or not the increase in the fraction of obscured star formation in high-redshift

clusters is related to infall. They speculate that the increase in star formation reflects the infall of new members and note that most of the MIPS-detected cluster galaxies are not projected onto the cluster core (inner 500pc). Over larger scales the work of Gallazzi et al. (2009) explored the obscured star formation fraction as a function of environment in the Abell 901/902 supercluster at  $z = 0.165$ . They find more obscured star formation at intermediate densities than in the cluster cores, similar to the distribution of the AGN population studied by Gilmour et al. (2007) in the same supercluster. If there is a substantial increase in the obscured star formation fraction in the intermediate densities around clusters, and the star formation in this environment increases with redshift, then projection of some of these structures onto the cluster core may contaminate the cluster estimates.

As discussed in Section 4.5, AGN in the large-scale environments around massive clusters may also project onto cluster cores. To better evaluate this possibility, it is useful to both directly measure the AGN population immediately outside clusters and measure the AGN population in intermediate densities more generally. Just as Poggianti et al. (2006) found that the fraction of [O II]-emitting galaxies increases in lower velocity dispersion environments, the AGN fraction as a function of environment is important because the environmental dependence may provide new information on the processes that drive AGN evolution. Both the XMM observations of the COSMOS fields (Silverman et al. 2009a,b) and *Chandra* observations of the Extended Groth Strip from DEEP2 (the All-wavelength Extended Groth strip International Survey, AEGIS; Georgakakis et al. 2008a; Georgakakis et al. 2008b) have estimated the AGN fraction in groups of galaxies or as a function of local overdensity at high redshifts. Georgakakis et al. (2008b) found that X-ray AGN are more frequently found in groups than in the field, which they connect to their observation that the X-ray AGN host galaxies are often red, luminous galaxies that tend to reside in denser environments, although they also find that this trend may reverse for the most powerful AGN. In a narrower redshift range from  $0.7 < z < 0.9$  and for  $M_B < -20$  mag they find that the AGN fraction is comparable in groups and the field and about 5%. This is approximately a factor of five higher than we find in clusters at similar redshifts, although these values are not exactly comparable as the Georgakakis et al. (2008b) AGN include somewhat lower-luminosity sources than our sample and the host galaxy magnitude limits are somewhat different. Silverman et al. (2009a) also investigate the environment dependence of X-ray AGN hosted by galaxies above a fixed stellar mass and find no strong preference between the field and groups except for the most massive galaxies, while Jeltema et al. (2007) find that the fraction of [O II]-emitting galaxies in intermediate-redshift, X-ray-selected groups ( $0.2 < z < 0.6$ ) is similar to clusters at the same redshift.

The clustering analysis by Coil et al. (2009) on the AEGIS data also helps to elucidate the distribution of AGN at high redshift as a function of environment, AGN luminosity, and host galaxy mass. They find that the X-ray AGN have similar clustering to luminous red galaxies and are more likely to reside in groups, while UV-bright QSOs are less strongly clustered and more similar to the field blue galaxy population. This is also similar to the results from Kauffmann et al. (2004) at low redshifts from SDSS, who find that galaxies at a fixed stellar mass that host luminous [O III] emission are twice as common in low-density regions as high. Taken together,

the AEGIS and COSMOS results illustrate that the measured AGN fraction depends on both the stellar mass (or luminosity) of the galaxy population and the star formation rate of the host, in addition to the AGN luminosity. This makes a direct comparison between these two surveys, as well as to our work on high-redshift clusters, somewhat problematic. The X-ray range considered by Silverman et al. (2009a) extends over  $42 < \log L_{0.5-10\text{keV}} < 43.7$ , or approximately half an order of magnitude below our X-ray threshold for a typical AGN SED. The X-ray AGN studied by Coil et al. (2009) extend an order of magnitude fainter than our work to a hard band limit of  $L_{X,H} > 10^{42} \text{ erg s}^{-1}$ . Both of these surveys are therefore dominated by intrinsically less luminous objects. The galaxy mass and luminosity ranges are similarly not identical. In future work we hope to put all of these high-redshift measurements on an equal basis for a more direct comparison.

While none of these results suggest that there are more luminous AGN in clusters than groups or the fields out to  $z \sim 1$ , such a trend may be seen at yet higher redshifts. Observations of cluster galaxies, particularly massive cluster ellipticals, suggest that most of their stars formed earlier than field galaxies (by 0.4 Gyr; van Dokkum & van der Marel 2007). If the central black holes of these galaxies grew contemporaneously, then perhaps by  $z \sim 2$  the AGN fraction will be higher in denser environments. Some interesting support for this picture comes from *Chandra* observations of the SSA22 protocluster at  $z = 3.09$  (Lehmer et al. 2009). They find a slightly higher AGN fraction in Lyman Break and Ly $\alpha$ -emitting galaxies in the protocluster compared to the field. While this is just one region, observations of the AGN fraction in clusters relative to the field at  $z \sim 2$  and above could provide interesting new insights into the coevolution of black holes and galaxies.

## 7. SUMMARY

We have conducted an expanded survey to identify luminous  $L_{X,H} \geq 10^{43} \text{ erg s}^{-1}$  AGN in clusters of galaxies from  $z \sim 0.05$  to  $z \sim 1.3$ . At low redshifts we have presented a new X-ray analysis of archival *Chandra* observations and spectroscopic follow-up of AGN candidates in six new clusters. There are no new, luminous AGN in these clusters and there are a total of just two luminous AGN in our sample of 17 clusters with  $z < 0.4$ . These measurements further strengthen the evidence for a very small luminous AGN fraction in low-redshift clusters. An important virtue of the new clusters is that the X-ray and spectroscopic coverage extends to the projected  $R_{200}$  radius and therefore they are better matched to observations of high-redshift clusters. At higher redshifts we have combined our previous work with literature data on X-ray sources, primarily from the ChaMP and SEXSI surveys, to compile a total sample of 15 clusters at  $z > 0.4$ . In spite of somewhat incomplete spectroscopic coverage of the X-ray sources in these fields, there are 18 luminous AGN in these clusters.

We parameterize the evolution of the AGN population in clusters in terms of the fraction of luminous galaxies that host AGN above our luminosity threshold. We have used a variety of techniques to estimate the number of luminous galaxies, defined to have  $M_R < M_R^* + 1$ , in these clusters and calculated the average cluster AGN fraction in several redshift bins. As the low and high-redshift clusters are reasonably well matched in terms of cluster velocity dispersion and X-ray temperature, the increase in the number of AGN is closely related to the increase in the fraction of galaxies more lumi-

nous than  $M_R^* + 1$ . Specifically, we find that the AGN fraction increases by approximately a factor of eight from  $z \sim 0.2$  to  $z \sim 1$ . This corresponds to an increase in the AGN population that scales as  $(1+z)^{5.3}$ . If the radio AGN population in clusters increases by a comparable amount, radio AGN may impact the identification of clusters as a function of redshift in current and planned SZ surveys. The substantial evolution in the cluster AGN population is also correlated with the evolution of the fraction of star-forming galaxies in clusters known as the Butcher-Oemler effect. Detailed studies of star formation and AGN in individual clusters could better quantify the extent that these two phenomena are coupled in clusters or perhaps even individual galaxies. We have also estimated the evolution of the field AGN fraction to compare it to the cluster AGN fraction. While the field AGN fraction is higher at all redshifts, the present data do not suffice to conclude if the rate of evolution is faster or slower in clusters. Future measurements of the relative evolution of star formation and black hole growth in clusters and the field could be an important probe of the coevolution of black holes and their host galaxies.

Measurements of the radial distribution of the cluster AGN provide new information on the origin of AGN within clusters. Unlike we found in previous work at low redshifts, the AGN in these high-redshift clusters are not strongly centrally concentrated when their distribution is plotted normalized to the  $R_{200}$  radius. This demonstrates that there are substantial numbers in the outskirts of clusters and supports the hypothesis that some cluster AGN are hosted by relatively gas-rich galaxies that have recently entered the cluster potential. While

this excess is not apparent in the velocity distribution, this may be due to biases in the measurement of the cluster velocity dispersion or simply small number statistics. We have also presented the first measurement of the XLF of cluster AGN at high-redshift and found that it is consistent with the field XLF at the same redshift. This comparison illustrates the future potential of XLF measurements in clusters to measure environment-dependent downsizing in clusters, as well as how the evolution of the cluster XLF can be used to constrain the evolution of black hole growth in clusters.

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